### HEATING OF ACTIVE REGION CORE LOOPS BY NANOFLARES



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# **Active Region Moss**



- Almost impossible to spatially resolve hot core loops.
- Moss regions are foot points of hot (3-5 MK) loops seen in core of active regions.

 Moss regions are bright, finely textured mottled, low lying emission above active region plage area, best seen in TRACE Fe IX/X 171 images (Schrijver et al 1999).





### Some results...

- Minimal variability in:
  - moss intensities (~10%) averaged over several TRACE pixels (Antiochos et al. 2003).
  - Doppler shift and line width (Brooks and Warren 2009).
  - density and temperature structure (Tripathi et al. 2010).
- Predicted intensities using steady heating match quite well with observed intensities IF an expansion at the footpoints of loops are considered (Warren et al. 2010).



#### Antiochos et al. 2003

Moss is time variable on a small scale. However global structure remains unchanged.

- These observations provide compelling evidence that heating is steady, but only if plasma does not have unresolved structures.
- Measurements suggest that there are unresolved plasma structures.
- We explore Emission Measure (EM) distribution in moss regions. Since these are the footpoints of the hot loops, the EM (T) is determined largely by variation of temperature along strands.
- We compare the **observed** EM(T) with **theoretical** EM(T) for different scenarios to gain some insight into the heating of hot core loops.

# **Transition Region DEM**

#### Klimchuk et al. (2008) derived DEM(T) for three limiting cases:

- Strong evaporation: the heat flux from the corona far exceeds the radiative losses from the transition region. The energy balance is then between thermal conduction heating and enthalpy cooling.
- Strong condensation: the heat flux from the corona much less than the radiative cooling. The energy balance is then between enthalpy heating and radiative cooling.
- Static Equilibrium: the heat flux from the corona very nearly balances the radiative losses from the transition region.

 $EM_{se} \approx ln(10) \left(\frac{\kappa_0}{14}\right)^{1/2} \frac{\bar{P} T^{3/4}}{k \Lambda(T)^{1/2}} \dots static equilibrium$ 

 $EM_{con} \approx -\ln(10) \frac{5 \ k \ J_0 \ T}{\Lambda(T)} \dots strong \ condensation$ 

$$EM_{ev} \approx \frac{ln(10)}{20} \frac{\kappa_0}{k^3} \frac{\bar{P}^2 T^{1/2}}{J_0} \dots strong \ evaporation$$

 $k_0 = 1.6 \times 10^{-6} \text{ CGS}$  k = Boltzmann's constant T = Temperature J0 = mass flux = n v n = electron number density v = plasma flow speed  $\Lambda(T) = \text{Optically thin radiative}$  loss function (from CHIANTI v6.0)

- Expressions provide emission measure curves for individual strands.
- We have considered average pressure and mass flux (J0) as arbitrary constants.
- We have restricted ourselves only with temperature dependance of the emission measure curves.







## Conclusions

- The EM(T) distribution for two different regions are strikingly similar.
- The EMs obtained using photospheric abundances are consistent for different ions and show a monotonically increasing trend from log T = 5.15 6.3.
- The EM for strong condensation best reproduce the observations supporting the idea that hot loops seen in the core of active regions are heated by nanoflares.

Tripathi, D., Mason, H.E., Klimchuk, J. A., 2010, to appear in The Astrophysical Journal