

The multi-thermal emission in solar active regions

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ABSTRACT

We present simultaneous SDO AIA and Hinode EIS observations of the hot cores of active regions (ARs) and assess the dominant contributions to the AIA EUV bands, extending our previous work. We find good agreement between SDO AIA, EVE and EIS observations, using our new EIS calibration and the latest EVE v.3 data. We find that all the AIA bands are multi-thermal, with the exception of the 171 and 335 Å, and provide ways to roughly estimate the main contributions directly from the AIA data. We present and discuss new atomic data for the AIA bands, showing that they are now sufficiently complete to obtain temperature information in the cores of ARs, with the exception of the 211 Å band. We found that the newly identified Fe XIV 93.61 Å line is the dominant contribution to the 94 Å band, whenever Fe XVIII is not present. Three methods to estimate the Fe XVIII emission in this band are presented, two using EIS and one directly from the AIA data. Fe XVIII emission is often present in the cores, but we found cases where it is formed at 3 MK and not 7 MK, the temperature of peak ion abundance in equilibrium. The best EIS lines for elemental abundance and DEM analysis are discussed. A new set of abundances for many elements are obtained from EIS observations of hot 3 MK loops. The abundances of the elements with low First Ionisation Potential (FIP), relative to those of the high-FIP elements, are found enhanced by about a factor of three, compared to the photospheric values. A measurement of the path length implies that the absolute abundances of the low-FIP elements are higher than the photospheric values by at least a factor of three. We present a new differential emission measure (DEM) method customised for the AIA bands, to study the thermal structure of ARs at 1'' resolution. It was tested on a few ARs, including one observed during the Hi-C rocket flight. We found excellent agreement between predicted and observed AIA count rates and EIS radiances. We found overall little differences in the AIA and Hi-C 193 Å images of coronal structures, despite the higher Hi-C resolution (0.25''). The Hi-C images and the AIA DEM modelling suggest that some of the cooler loops (below 1 MK) are already resolved by AIA, while the hotter (1.5–2.5 MK) ‘background’ emission is in most places still unresolved even at the Hi-C resolution. This unresolved emission is significantly lower than previously observed with TRACE and the SOHO CDS and Hinode EIS spectrometers. Its enhancement appears to be mostly due to increased iron abundance. We find an ubiquitous presence of emission at different temperatures that is not co-spatial, and suggest that future high-resolution imaging is done with isothermal bands.

Key words. Atomic data – Line: identification – Sun: corona – Techniques: spectroscopic

1. Introduction

Understanding the thermal structure of the coronal emission in solar active regions (ARs) is a key element for our understanding of the physical mechanisms responsible for their heating. Observational and theoretical aspects have been debated for a very long time. For recent reviews, see e.g. Klimchuk (2010) and Reale (2010). This paper focuses on a few observational aspects, combining the latest imaging and spectroscopic observations with atomic data and modelling.

Most of the observational facts about active regions have been known from the 1970s thanks to the Skylab observations. Key observations were made even before though, from the ground. Many important aspects have been found from later missions such as SMM, Yohkoh, SOHO, TRACE, Hinode. The SOHO Coronal Diagnostic Spectrometer (CDS) provided since 1996, for the first time routinely, spatially resolved monochromatic images of active regions over a wide temperature range. The instrument was carefully calibrated in-flight (see Del Zanna et al. 2010a). The Hinode EUV Imaging Spectrometer (EIS) with its two channels (SW: 166–212 Å; LW: 245–291 Å) has provided since 2006 AR observations in a wide range of temperatures (from e.g. all the iron ionisation stages from Fe VII to Fe XVI) and above 10 MK (from Fe XXIII and Fe XXIV, as presented in detail in Del Zanna 2009a, 2012b, 2008; Del Zanna et al.

2011b). We have also recently provided an in-flight calibration for EIS (Del Zanna 2013). These spectrometers have great temperature resolution but poor spatial (CDS: 4–5''; EIS: 3–4'') and temporal resolution.

Imaging from TRACE had excellent spatial resolution (about 1'') but limited temperature resolution, and significant stray light. The Atmospheric Imaging Assembly (AIA, see Lemen et al. 2012) on-board the Solar Dynamic Observatory (SDO), launched in 2010, has been observing, for the first time routinely, the solar corona at high resolution (about 1'') and 12s cadence with six EUV bands. These observations allow in principle the study of the thermal structure of active regions at high spatial resolution. Another significant advance compared to previous instruments is the low stray light, which has been recently assessed by Poduval et al. (2013).

The highest-resolution images (about 0.25'') of the solar corona have recently been taken with an excellent rocket flight carrying a telescope with a broad band similar to the AIA 193 Å one (Cirtain et al. 2013). These images provide important information on the small-scale structures in ARs.

Active regions are notoriously bright in the X-rays in their cores. The Skylab X-ray imaging suggested that the long-lasting ‘hot’ core loops in quiescent ARs are approximately isothermal around 3 MK (Rosner et al. 1978), a puzzling feature. Later spectroscopic measurements in the X-rays from SMM/FCS con-

firmed that most active region cores were nearly isothermal, at least in the sense that very little material at temperatures above 3 MK is present (see, e.g. Schmelz et al. 1996). EUV spectroscopy with SOHO CDS also confirmed this (Del Zanna & Mason 2003). For simplicity, we will call these hot core loops ‘3 MK’ loops, bearing in mind that the isothermality of these 3 MK loops is not universally accepted. Various recent studies based on Hinode EIS observations discussed AR cores (e.g. Tripathi et al. 2011; Warren et al. 2012) but did not analyse in detail the thermal structure of the 3 MK loops. Winebarger et al. (2011) calculated the differential emission measure (DEM) of the apex region of the hot core loops, and found the DEM to be broad and peaked around 3 MK.

We note that previous studies used the Hinode EIS ground calibration, hence need to be revised in light of our new calibration (Del Zanna 2013), which mainly increases the lines of the LW channel. This affects by about a factor of two the emission measures at 3 MK, since they are constrained by the strong Fe XIV–Fe XVI observed in this channel. We also note that, in general, it is not simple to determine how much isothermal a plasma is, with significant uncertainties originating from how accurate observations and atomic data are (see, e.g. Judge 2010; Landi & Klimchuk 2010).

One of the aims of this paper is to study the thermal structure of these 3 MK loops taking into account the new calibration, a different selection of lines, a discussion of elemental abundances, and background/foreground subtraction, all aspects that were often neglected in the literature but that turn out to be quite important. These loops appear ‘fuzzy’, i.e. unresolved at 1'' resolution. This was known from Skylab X-ray imaging, EUV imaging from TRACE (see, e.g. Brickhouse & Schmelz 2006, and actually for a long time from ground-based observations (e.g. in Ca XV), although detailed measurements of densities and abundance have been lacking).

In the EUV, the brightest emission in ARs comes from low-lying emission (commonly known as ‘moss’) that was clearly visible with high-resolution imagers observing 1 MK emission such as NIXT (Peres et al. 1994) and TRACE 173 Å passband (see, e.g. Schrijver et al. 1999). Martens et al. (2000) proposed that moss is the transition-region emission of the 3 MK loops, a view commonly accepted. We note however that moss emission is co-spatial with all plage regions, and plage regions are the magnetic footpoint of all coronal structures, i.e. not only the 3 MK loops. Various studies on moss emission have been performed, to try and gain insight into the physics of the 3 MK loops (see, e.g. Tripathi et al. 2010; Winebarger et al. 2011; Tripathi et al. 2012). We will discuss moss characteristics in a separate paper, but note that any measurement of coronal structures needs to avoid regions where moss is present.

Active regions, especially during later stages of their lives, produce long-lasting ‘warm’ loops, formed between 1 and 2 MK, spectacularly observed with the TRACE 173 Å (Fe IX) band. The thermal characteristics of these loops have been debated for a very long time. Clearly, EUV imaging from TRACE had limited temperature information. Until 2003, it was commonly thought that the majority of AR loop structures were multi-thermal, and that most cooler loops were actually hot loops observed during the cooling process.

Our SOHO CDS observations, combined with TRACE high-resolution imaging of many ARs actually showed a very different picture. Most warm loops were found nearly isothermal in their cross-section and with filling factors close to one at their base (Del Zanna 2003a; Del Zanna & Mason 2003; Certain

et al. 2007). Our results have so far been largely confirmed by Hinode EIS and SDO AIA observations. For a few examples, see e.g. Warren et al. (2008); Tripathi et al. (2009); Aschwanden & Boerner (2011); Schmelz et al. (2013), although some loops show some departures (Schmelz et al. 2009, 2010), and the issue is still debated.

In summary, our view is that a significant fraction of the coronal emission in active regions is composed of loop structures that are nearly isothermal in their cross-section, hence are better observed with spectrometers with high spatial resolution and/or with monochromatic imagers. Clearly, given the poor spatial resolution of current spectrometers, most loop structures appear as multi-thermal, due to line of sight effects. It is thought that individual magnetic flux tubes are isothermal, and that resolved loops are composed of a superposition of such elementary structures (Klimchuk 2006). Measuring how isothermal observed loops are, is therefore an useful observational constraint for loop modelling.

The effective spatial resolution of EIS is only marginally better than CDS. The AIA resolution is excellent, but one obvious question we want to address here is whether the nature, the calibration and the atomic data for the AIA EUV bands are accurate enough to measure the thermal structure of active regions.

As we showed in O’Dwyer et al. (2010) and Del Zanna et al. (2011c), the AIA EUV bands are significantly multi-thermal, which naturally ‘blurs’ their view of active region loops and limits their use. The O’Dwyer et al. (2010) study was purely theoretical, while the Del Zanna et al. (2011c) study combined Hinode EIS and AIA observations to show that significant ‘cool’ ($T < 1$ MK) emission is present in all the AIA EUV bands. EIS has a significant temperature overlap with the AIA bands, and further observes directly some of the AIA bands, so it is an excellent instrument for this purpose. We highlighted several problems in the AIA and EIS data, but also in the atomic data, especially for the 211 Å band, where up to a half of the emission can come from unidentified spectral lines. In this paper we aim to extend the previous work to discuss the various contributions to the EUV AIA bands in the case of hot emission (cores of active regions), and to see how the AIA data can be used to study the AR multi-thermal structure with differential emission measure (DEM) modelling.

Our SOHO CDS observations also suggested the presence of a pervasive and dominant unresolved emission in the 1.5–2.5 MK temperature range (Del Zanna & Mason 2003), which constituted about 80% of the EUV AR emission observed by CDS and TRACE, and was neglected in previous analyses. Similar levels of emission are also observed with EIS, but not with AIA as we shall discuss here. Another natural question we want to address here is whether this unresolved emission is real or is an artifact of the low spatial resolution of the spectrometers and stray light in the imagers.

2. AIA and EIS Observations

We have analysed several SDO/AIA and Hinode/EIS observations, both of on-disk and off-limb active regions. We used an EIS study (Atlas_60) that we designed for the purpose of the cross-calibration between AIA and EIS. Here, we present a sample of results from two on-disk active regions, observed on 2010 Nov 23 and 2010 Oct 26. The first one was discussed in Del Zanna et al. (2011c) as a clear example of a case when significant cool emission is present in all the AIA EUV bands. It is also a clear case where hotter emission from Fe XVIII is not present in

the 94 Å band. The second one is instead a clear case where Fe XVIII is present in the 94 Å band. We also analysed with the same level of detail an AR observation of 2010 Dec 11 at the limb, where some Fe XVIII emission is present in the core.

The EIS observations consists of single rastering of the 2'' slit (from west to east) 60 times by $\approx 2''$ steps, for a field-of-view (FOV) of $120'' \times 140''$. With 60 s exposure time, the observation lasted ≈ 60 minutes. Details on the complex processing of the EIS and AIA data can be found in Del Zanna et al. (2011c). Various geometrical corrections were applied to both AIA and EIS. The AIA plate scale was adjusted for each observation by examining the solar limb, and then cross-checked by cross-correlating low-lying features such as the moss. The co-alignment is good to one AIA pixel. The AIA images were then reduced to the lower spatial and temporal resolution of the EIS data for a direct comparison. The AIA images were convolved with a Gaussian of 2'' full-width-half-maximum (FWHM). Various comparisons on different ARs were performed, and the results consistently indicate an effective resolution of EIS of about 3–4''. This is due to a superposition of various effects which include the intrinsic EIS point spread function, its stray light, and the instrument and satellite jittering during each exposure. The long exposures required for a spectrometer, combined with the intrinsic AR variability, also contribute to ‘blurring’ the view.

3. Atomic data and instrument calibrations

In the past few years, we (the APAP team¹) have produced a significant amount of new large-scale atomic calculations for a large number of ions, including Fe VII (Witthoeft & Badnell 2008), Fe VIII and Fe IX (O’Dwyer et al. 2012), Fe X (Del Zanna et al. 2012b), Fe XI (Del Zanna & Storey 2013a), Fe XII (Del Zanna et al. 2012a), Fe XIII (Del Zanna & Storey 2013b), Fe XIV (Liang et al. 2010), and Fe XVI (Liang et al. 2009). The R-matrix calculations for Fe X–Fe XIII were not trivial, and a lot of surprising issues were found. The main being that cascading and resonance excitations are extremely important for some of the soft X-ray and EUV lines. These atomic data have been benchmarked against the best solar and laboratory observations in a series of papers, where several new identifications have been provided. The review of the soft X-rays is given in Del Zanna (2012a). For the present analysis we adopt the above atomic data and line identifications, and CHIANTI v.7 for the other ions. Note that some of the new data and identifications have already been made available to CHIANTI v.7.1 (the Fe VIII, Fe IX and Fe XIV atomic data, and some of the Del Zanna 2012a identifications). The latest v.7.1 CHIANTI ionization equilibrium tables (Landi et al. 2013) are adopted here. A notable feature is a change in the ionization equilibrium for Fe VIII and Fe IX, compared to the previous one (Dere et al. 2009).

During the present analysis, significant discrepancies in the radiances of the EIS lines were found, with most lines in the long-wavelength (LW) channel being a factor of two too weak, compared to those in the short-wavelength (SW) band, if the ground calibration is adopted. A detailed analysis has led to a new time- and wavelength-dependent radiometric calibration (Del Zanna 2013) which is adopted here. The new calibration increases all the LW lines by about a factor of 2 for the observation discussed here, but leaves those of the SW band substantially unchanged (in agreement with the in-flight calibration discussed in Wang et al. 2011).

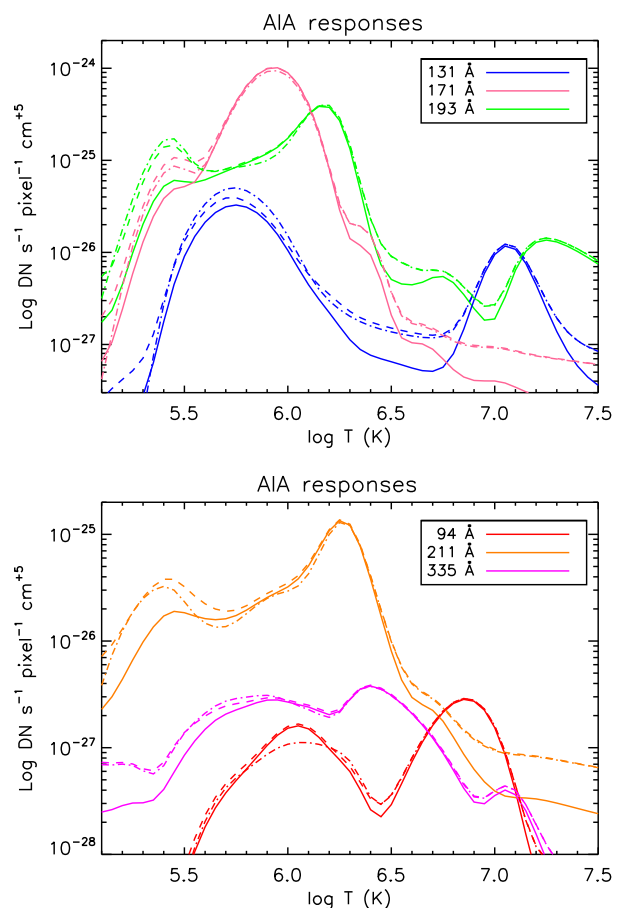


Fig. 1. The SDO AIA response functions calculated with the present atomic data, the present ‘coronal’ abundances and a constant electron density of $2 \times 10^9 \text{ cm}^{-3}$ (full lines). The dashed lines are the responses calculated with ‘photospheric’ abundances, multiplied by 3.16. The dash dot lines are calculated with CHIANTI v.7.1, and ‘photospheric’ abundances multiplied by 3.16.

An assessment of the AIA calibration for data in 2010 Dec 10 has been done by comparing the SDO Extreme ultraviolet Variability Experiment (EVE) (Woods et al. 2012) and AIA observations. We compared the full-disk AIA data numbers (DN) during the first hour of 2010 Dec 10 with the values estimated directly from EVE, using the AIA pre-flight calibration of Boerner et al. (2011). We measured the radial distribution of the off-limb AIA count rates and estimated the missing AIA off-limb contribution to be negligible (of the order of 1%). With the EVE v.2 data, we obtained the following correction factors to the AIA effective areas: 0.79, 0.84, 1.28, 1.13, 1.12 for the 131, 171, 193, 211, and 335 Å bands. The correction factors suggested within the AIA software (keywords `evenorm` and `timedepend_date`) are very similar (0.78, 0.83, 1.24, 1.09, 1.08). However, with the recent EVE v.3 calibration, we obtained different values: 0.83, 1.11, 1.12, 0.97, 1.0. These corrections have not been applied, since they are well within the uncertainties of the EVE calibration (25%). We also note that a proper in-flight AIA calibration is non-trivial. For example, any correction to the 94 and 171 Å bands has an additional uncertainty due to the low EVE resolution.

We used the present atomic data and the pre-flight effective areas (Boerner et al. 2011) to calculate the AIA response functions, following the simple methods described in Del Zanna

¹ www.apap-network.org

et al. (2011c). These responses are basically the sum of the emissivities of the lines and continua folded with the effective areas and summed over wavelength. We calculated the responses at a constant electron density of $2 \times 10^9 \text{ cm}^{-3}$, typical for the upper transition-region emission in warm loops (Del Zanna & Mason 2003), but note that some variations would be present if different densities are used. Note that the AIA response functions available within Solarsoft assume a constant pressure of $10^{15} \text{ cm}^{-3} \text{ K}$ and a set of ‘coronal’ elemental abundances.

We adopted a set of ‘coronal’ abundances discussed below, where most low first ionization potential (FIP) elements such as iron have been increased by a factor of 3.16. The curves are shown in Fig. 1. We also calculated the responses with the photospheric abundances of Asplund et al. (2009). These values, increased by 3.16, are also shown in the figure. The peaks agree very well because the dominating lines in all the AIA EUV bands are from iron, so the abundances are mainly a scaling factor, although some differences are present in the continuum. This means that emission measure modelling with AIA is largely independent of elemental abundances. Finally, the dot-dash lines in Fig. 1 show the responses calculated with CHIANTI v.7.1.

3.1. AIA 131 Å

The AIA 131 Å band is dominated by two Fe VIII lines at 130.94 and 131.24 Å although significant contribution in ARs comes from a large number of lines, plus the free-bound continuum. The atomic data for Fe VIII are uncertain. As described in Del Zanna (2009b), adjustments to the Griffin et al. (2000) atomic data were necessary to improve the EUV lines observed by Hinode EIS. These adjustments, implemented in CHIANTI v.7, were based on a large atomic structure calculation that focused on the lower levels and was not accurate for the 4f levels that are producing the lines in the AIA 131 Å band. A recent large-scale atomic calculation from Tayal & Zatsarinny (2011) did not resolve the problems with the EUV lines (cf. their Table 4), but should in principle provide better atomic data for the 4f lines. We therefore built a new ion model where we used the Tayal & Zatsarinny (2011) data but also applied some corrections, based on the Del Zanna (2009b) calculations for the levels numbered 41, 43, 45, and 46. Tayal & Zatsarinny (2011) only provided some A values for the dipole-allowed transitions, however a significant number of forbidden transitions are necessary to redistribute the level population, especially to de-populate the metastable levels for this ion. The values calculated by Del Zanna (2009b) have been added to those provided by Tayal & Zatsarinny (2011). The new ion model produces a significant (about 40%) decrease in the 131 Å band response, as shown in Fig. 1.

When flare-like emission is present, significant continuum and Fe XXI (formed around 10 MK) emission is present in the 131 Å band, with some contribution also from Fe XX and Fe XXIII (see Petkaki et al. 2012; Del Zanna & Woods 2013). This hot contribution significantly complicates the use of this band for DEM analysis, because it is typically superimposed on the cooler (mostly Fe VIII) emission, for which we have no other means to establish its amount. Looking at many ARs, we found that only occasionally the Fe XXI is produced in the same regions as Fe XVIII, observed in the 94 Å band (see below). This means that the estimated Fe XVIII cannot be used as a proxy to estimate the Fe XXI emission. Also, the DEM that can be estimated from Fe XVIII is not necessarily the same as that for Fe XXI. So, as a first approximation, we did not use the 131 Å band for the DEM

analysis for all the regions where the estimated Fe XVIII emission is above a threshold value. In case of obvious flaring and Fe XXI contribution, the 131 Å band is used for the DEM analysis.

3.2. AIA 171 Å

The 171 Å band turns out to be always dominated by the resonance transition from Fe IX, for which the atomic data of Storey et al. (2002) are expected to be reliable. This line is predicted to be sensitive to both the temperature and density structure of the emitting plasma. As shown in Del Zanna et al. (2011c), this line is formed over a very broad temperature range ($\log T[\text{K}] = 5.5\text{--}6.2$). In the cores of active regions, some blending (about 10% of the Fe IX) is due to Ca XV 171.6 Å.

The 171 Å line is barely visible in the EIS spectra, being in a region where the EIS sensitivity is about three orders of magnitude lower than at the peak around 195 Å. A direct AIA vs. EIS comparison is therefore difficult.

3.3. AIA and Hi-C 193 Å

This band is relatively well understood, because all the lines are observed with EIS, and since 2004 we have made a significant effort to identify the lines and provide accurate atomic data especially for Fe X, Fe XI, and Fe XII. We note that the strongest lines in this band are from Fe XII. When flare-like emission is present in the cores of ARs, the Ca XVII 192.8 Å and Fe XXIV 192.0 Å lines provide a significant contribution to this band. Near the footpoints of warm loops, Fe VII, Fe VIII, and Fe IX produce a significant contribution (Del Zanna et al. 2011c). We note that Fe VII is not included in CHIANTI yet because of uncertainties in the identifications of the strongest lines. Here, we adopt the identifications of Del Zanna (2009a). Some unidentified lines are also present in this band (Del Zanna et al. 2011c), although they contribute little in most AR areas.

Despite common knowledge, only about half of the counts recorded in the AIA 193 Å band in the cores of active regions is due to the Fe XII lines. The rest of the counts are due to several transitions formed over a broad range of temperatures. There is significant contribution from Fe XI, then from Fe X and Fe XIII. We have found this by comparing the simultaneous AIA and EIS observations of a few ARs. Fig. 2 shows as an example the AR observed on-disk on 2010 Nov 23. Its top right image shows the total AIA count rates due to the Fe XII 186.8, 192.4, 193.5, and 195.1 Å lines, all observed by EIS. These count rates were obtained by converting the EIS calibrated radiances into AIA DN/s. They are to be compared to the total AIA count rates, shown in the top left image.

By comparing the AIA count rates due to Fe XII (as estimated by EIS) with the AIA images, we found that the DN/s due to all the Fe XII lines in the 193 Å band can be roughly estimated in this way:

$$I(\text{Fe XII}) = 0.7 \times (I(193 \text{ \AA}) - I(171 \text{ \AA})/2)$$

Fig. 2 shows (middle top) an image of the AIA DN/s due to the Fe XII lines as obtained from this expression. A comparison with the measured Fe XII intensity shows that this linear relation overestimates the Fe XII contribution in the areas with strong cool emission and little Fe XII emission, as for example pixel position 25 in the lower panel of Fig. 2.

We note that for the present DEM analysis the total observed and predicted emission in this band has been used, since the

cooler contribution to this band can be estimated quite accurately with the 131 and 171 Å bands, and the atomic data for the 193 Å band are relatively complete.

3.4. AIA 211 Å

This AIA band is the least understood, with a significant number of lines not yet identified, as discussed in Del Zanna et al. (2011c). Some of them are ‘cool’ lines (Del Zanna 2009a), while others are coronal (Del Zanna 2012b). The main coronal line is the Fe XIV 211.3 Å transition. This Fe XIV line is observed directly by Hinode/EIS, so we can estimate precisely its contribution to the AIA band (assuming that the EIS and AIA calibrations are correct). Fig. 2 (right middle image) shows the AIA DN/s due to the Fe XIV 211.3 Å line. In the cores of ARs, only about 60% of the counts in the 211 Å band comes from the Fe XIV transition at 211.3 Å.

The missing atomic data for this band means that it cannot be used as it is for *DEM* analysis. In the absence of EIS observations, the AIA 211 Å DN/s due to the Fe XIV line can be roughly estimated by subtracting a linear combination of the 171 and 193 Å bands

$$I(\text{Fe XIV}) = I(211 \text{ \AA}) - I(171 \text{ \AA})/17 - I(193 \text{ \AA})/5$$

The result obtained from this relation is shown in Fig. 2 (middle image). We have used this linear combination as input for all the *DEM* analyses presented in this paper. The corresponding response for this band was calculated by adding the Fe XIV emissivity to the continuum. The atomic data for Fe XIV are reliable (Liang et al. 2010; Del Zanna 2012b) and the 211.3 Å line does not have significant temperature or density effects.

3.5. AIA 335 Å

The 335 Å band is also highly multi-thermal, although the dominant contribution in the cores of active regions is from a single Fe XVI 335.4 Å transition, for which the atomic data (Liang et al. 2009) are reliable. We have checked this by estimating the Fe XVI 335.4 Å line using the intensity of the Fe XVI 263 Å line observed by EIS, and adopting a theoretical ratio of 17.5 (photons) for the two lines. Fig. 2 (bottom right image) shows the estimated AIA 335 Å DN/s due to the Fe XVI 335.4 Å line. The residual emission in AR cores comes from a host of transitions and ions (Fe XIV, Mg VIII, Fe XII, etc.). We found that the empirical relation

$$I(\text{Fe XVI}) = I(335 \text{ \AA}) - I(171 \text{ \AA})/70 ,$$

provides a rough estimate for the AIA DN/s due to the Fe XVI 335.4 Å line in the AR cores. Fig. 2 (middle lower image) shows the result of this linear relation. This empirical adjustment sometimes leaves a small residual due to the cooler lines in regions where the Fe XVI 335.4 Å line is very weak.

We have used this estimate as input in some *DEM* analyses, using for the response function just the Fe XVI 335.4 Å contribution function, added to the continuum. This procedure effectively removes the double peak for this band. The lower temperature peak in this band is due to significant TR emission, as discussed in Del Zanna et al. (2011c) (see the cool loops in the bottom left of the images in Fig. 2).

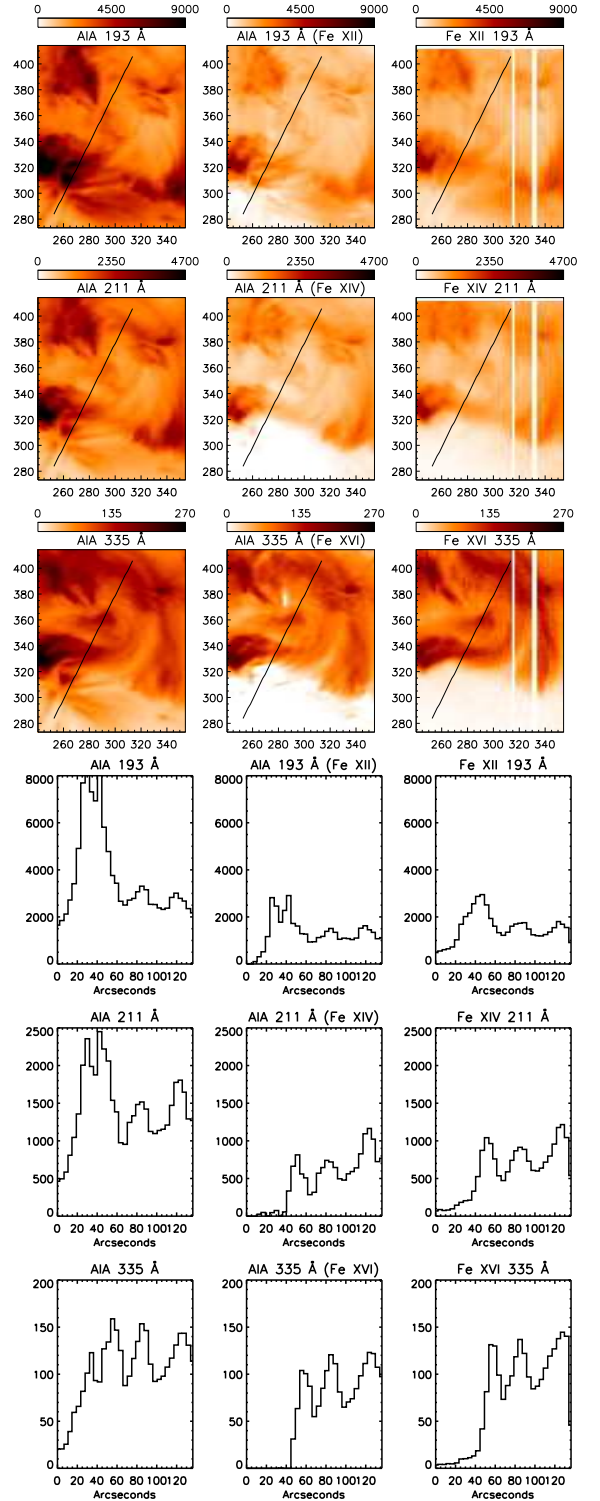


Fig. 2. Top three left images: AIA count rates (DN/s) of the on-disk AR observation of 2010 Nov 23, rebinned to the EIS resolution. Top three middle images: estimated AIA count rates due to Fe XII, Fe XIV, and Fe XVI (see text). Top three right images: measured Fe XII, Fe XIV and estimated Fe XVI (in AIA DN/s) directly from EIS. The axes in the images indicate arc seconds from Sun centre. The bottom panels show the corresponding profiles (AIA DN/s) along the segment indicated in the images. The AR core spans the 60–130 arc seconds region in the profiles.

3.6. AIA 94 Å

This band is particularly important for the presence of the Fe XVIII 93.932 Å line. Significant problems in the 94 Å band

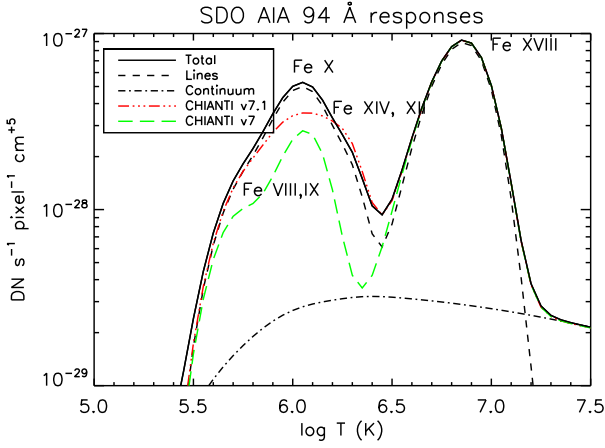


Fig. 3. AIA 94 Å responses

were reported by several authors (see, e.g. Aschwanden & Boerner 2011; Del Zanna et al. 2011c; Schmelz et al. 2011; Foster & Testa 2011; Testa et al. 2012b), but it was not clear if they were due to calibration or atomic data issues, so a few authors produced empirical estimates for the Fe XVIII emission (Reale et al. 2011; Testa & Reale 2012; Warren et al. 2012).

When the Fe XVIII emission becomes dominant, we found good agreement between observations and theory (Petkaki et al. 2012), suggesting that the AIA 94 Å calibration is correct, given that the atomic data we provided for Fe XVIII should be reliable (Witthoef et al. 2006; Del Zanna 2006).

Outside of obvious flaring emission, the 94 Å band was thought to be dominated by Fe X 94.012 Å, identified by Edlén (1937). Del Zanna et al. (2011c) however pointed out that at least two other unidentified components were present in ARs, one ‘cooler’ below 1 MK, and one ‘hotter’ around 2–3 MK.

The main problem for this band was the lack of atomic data and line identifications in the soft X-rays. As already mentioned, we have performed new large-scale R-matrix calculations for several iron ions, to provide the data for the soft X-ray lines. The soft X-rays (which include this band) were particularly problematic because the majority of lines were unidentified. These new atomic data have now been benchmarked against soft X-ray solar and laboratory observations in Del Zanna (2012a), where several new identifications (for all the strongest lines) were provided. Of all these new identifications, the Fe XIV 93.61 Å transition is of particular importance for AR observations in the AIA 94 Å band, because it becomes very strong. Indeed, it has been known for a very long time (see, e.g. Widing & Sandlin 1968; Manson 1972) that when the Sun is active (because of the presence of non-flaring ARs) the irradiance of the 93.61 Å line becomes almost as strong (66%) as that of the Fe X 94.012 Å transition. Clearly, this line must become stronger than Fe X in AR radiance observations. Note that the atomic data for Fe XIV (Liang et al. 2010) are reliable, and that the new identifications has been made available to version 7.1 of the CHIANTI database (Landi et al. 2013).

The Fe X $3s^2 3p^5 \ ^2P_{3/2} - 3s^2 3p^4 4s \ ^2D_{5/2}$ 94.012 Å transition, identified by Edlén (1937), is the dominant emission in quiet Sun conditions. The atomic calculations for the Fe X $3s^2 3p^4 4s$ levels were extremely complex, as discussed in Del Zanna et al. (2012b). Our R-matrix large-scale calculations should be accurate to within 30%. Distorted wave (DW) calculations underestimate significantly the Fe X lines. CHIANTI version 7.1 (Landi et al. 2013) included DW data (Landi & Dere 2013). These lat-

ter calculations predict the presence of a $3s^2 3p^5 \ ^2P_{3/2} - 3p^6 3d \ ^2D_{5/2}$ transition, at an estimated wavelength of 94.27 Å. There are two problems with this line. First, its intensity is predicted to be more than half the 94.012 Å line, while there are no lines in astrophysical or laboratory spectra around that wavelength. Second, the position of the $3p^6 3d \ ^2D_{5/2}$ level affects the $3s^2 3p^4 4s \ ^2D_{5/2}$ one, drastically reducing the weighted oscillator strength of the 94.012 Å transition to a value of 0.13, compared to a value of 0.31 obtained from the large-scale RM4 calculation presented in Del Zanna et al. (2012b). The benchmark of the Fe X soft X-ray lines presented in Del Zanna (2012a) shows very good agreement between the RM4 model and observations, suggesting problems in the target used in Landi & Dere (2013).

Significant contribution from the continuum (mostly due to free-bound emission at 3 MK) is present in the AIA 94 Å band (as discussed below), although its amount depends on the absolute abundances.

Weak Fe VIII and Fe IX transitions are also present in this band. The Fe VIII lines were observed in the laboratory and DW data were calculated (O’Dwyer et al. 2012) and included in CHIANTI version 7.1. DW calculations for the Fe IX lines (O’Dwyer et al. 2012) were also included in CHIANTI version 7.1, however the wavelengths of these lines are very uncertain. The benchmark work of Del Zanna (2012a) found problems in the suggestions given in O’Dwyer et al. (2012) (included in CHIANTI version 7.1.), providing new recommended values.

CHIANTI version 7.1 also predicts a few Fe XII lines that could contribute to the 94 Å AIA band, in particular a $3s 3p^4 \ ^2D_{5/2} - 3s^2 3p^2 4p \ ^2D_{5/2}$ transition. In this case, both the DW calculations of Landi & Dere (2013) and the large-scale R-matrix one (Del Zanna et al. 2012a) agree. The estimated wavelength in CHIANTI version 7.1 is 93.68 Å, while that one estimated by Del Zanna (2012a) is 93.49 Å. Note, however, that this wavelength has a significant uncertainty, considering that large discrepancies in the energies of the $3s^2 3p^2 4p$ levels have been found (Del Zanna 2012a).

Fig. 3 shows the AIA 94 Å responses calculated with the present data, together with previous CHIANTI versions. The increase around $\log T[\text{K}] = 6.3$ is mainly due to the Fe XIV 93.61 Å line. The main difference with CHIANTI v.7.1 is due to different atomic data for Fe X. We note that various comparisons with observations suggest that the new atomic data (which will be available in version 8 of CHIANTI) are relatively accurate and complete. The use of semi-empirical fudge factors such as the `/chiantifix` keyword within the AIA software is discouraged.

In conclusion, the AIA 94 Å band is so multi-thermal that it is not very useful for DEM analysis as it is. However, we have used it to constrain the DEM at higher temperatures, when significant Fe XVIII is present. We searched for a linear combination of the 171 and 211 Å bands that would produce similar count rates as those we estimated to be due to Fe XVIII (as described below). We found that a rough estimate of the AIA 94 Å DN/s due to Fe XVIII is

$$I(\text{Fe XVIII}) = I(94 \text{ \AA}) - I(211 \text{ \AA})/120. - I(171 \text{ \AA})/450.$$

The 211 Å emission is used as a proxy for the Fe XIV (recall that the Fe XIV 211.3 Å transition produces about half of the count rates) and the cooler (1 MK) emission. The 171 Å emission is instead used to roughly estimate the weaker contributions from Fe VIII, Fe IX, and Fe X lines.

The estimated AIA Fe XVIII count rates are similar to those obtained using the Warren et al. (2012) 171, 193 Å combination

(note that the sum in the paper should be from 0 to 3), as shown below. Note that the Warren et al. (2012) empirical combination produced a good agreement with the observed Fe xviii emission (Teriaca et al. 2012) for one sample AR. The agreement between the present and the Warren et al. (2012) method is not surprising, considering that the morphology of the 193 and 211 Å bands is very similar in all active region cores.

4. DEM inversion methods

We recall that the intensity of an optically thin line can be expressed as an integral along the line of sight

$$I(\lambda_{ji}) = \int_h N_e N_H Ab G(N_e, T, \lambda_{ji}) dh \quad (1)$$

where Ab is the elemental abundance, $G(N_e, T, \lambda_{ji})$ is the *contribution function* of the spectral line. $G(N, T)$ has a very strong dependence on temperature, through the ionization balance calculations, and normally a negligible dependence on the electron number density, N_e , for a certain range of densities, if the appropriate line is chosen. In the case of a broad-band instrument such as AIA, the contribution function is replaced by the instrument response, which is calculated by summing all the $G(N, T)$ of the lines in the passband and adding the emissivity of the continuum.

If a unique relationship exists between N_e and T , neglecting density effects we define

$$DEM(T) = N_e N_H \frac{dh}{dT} \quad [cm^{-5} K^{-1}] \quad (2)$$

as the column differential emission measure (*DEM*) of the plasma, which gives an indication of the amount of plasma along the line of sight that is emitting the radiation observed at a temperature between T and $T + dT$. The *DEM* by definition is a continuous distribution in a range of temperatures. We normally write

$$I(\lambda_{ji}) = Ab \int_T G(T) DEM(T) dT, \quad (3)$$

assuming that the abundance of the element Ab is constant along the line of sight. There is an extended literature on the subject of inversion techniques and the non-uniqueness of the solution (see e.g. Craig & Brown 1976, 1986; Judge et al. 1997 and references therein). A few specific issues are discussed in Del Zanna (1999); Del Zanna et al. (2002).

The column emission measure $EM_h(T, \Delta T)$ can be calculated by integrating the *DEM* over the temperature bins $\Delta T = T_2 - T_1$

$$EM_h(T, \Delta T) \equiv \int_{h1}^{h2} N_e N_H dh = \int_{T1}^{T2} DEM(T) dT \quad [cm^{-5}] \quad (4)$$

The total emission measure $EM_h(T)$ is obtained by integrating over the entire temperature range.

Note that the absolute values of the *DEM* or EM_h depend on the radiometric calibration of the instrument providing the intensities and the ‘absolute’ (i.e. relative to hydrogen) values of the elemental abundances. Also note that there is a vast literature on the subject of photospheric and coronal abundances. Only a few key issues are mentioned here. The standard photospheric abundances of many elements have recently changed significantly. The low-FIP (≤ 10 eV) elements are more abundant than high-FIP in the slow solar wind (SSW) and in coronal loop plasma.

Their relative values compared to hydrogen is still debated, as well as the amplitude of the difference. For recent reviews see, e.g. Asplund et al. (2009); Lodders et al. (2009); Schmelz et al. (2012). A common set of ‘coronal’ abundances used by many authors is that one of Feldman et al. (1992), where the low-FIP elements are increased by a factor of about 4. This set of abundances is in need of a significant revision, mainly because many values were measured relative to the photospheric abundances of high-FIP elements (such as oxygen), which have changed. Also, we note that it is now quite clear that a single set of elemental abundances cannot be representative of all the features of the solar corona. For example, the quiet Sun corona shows pretty much photospheric abundances. A strong FIP effect was thought to be present in warm loops, however a revision of the older measurements, together with SOHO CDS measurements has showed that in many cases warm loops have abundances close to the photospheric ones, i.e. the FIP enhancement was often overestimated in the literature (see Del Zanna 2003b). We discuss the abundances in hot core loops below.

An interesting measurement is obtained once the column emission measure and the average electron density $\langle N_e \rangle$ are known. This is the path length

$$\Delta h = \frac{EM_h}{0.83 \langle N_e \rangle^2} \quad [cm] \quad (5)$$

i.e. the length along the line of sight of the emitting plasma, assuming that it is uniformly distributed. 0.83 is the value of N_H/N_e (which depends slightly on the elemental abundances). The average electron density $\langle N_e \rangle$ is obtained from density-sensitive line ratios, is independent of the elemental abundances, and largely independent of the ionisation state of the plasma. Another parameter often used is the spectroscopic filling factor, which is obtained from the path length once the size (along the line of sight) of the plasma is known. This size can be estimated quite accurately only in some cases, such as clearly defined coronal loops.

One useful parameter related to the *DEM* is the effective temperature

$$T_{\text{eff}} = \int G(T) DEM(T) T dT / \left(\int G(T) DEM(T) dT \right) \quad (6)$$

that is an average temperature more indicative of where a line is formed. This is often quite different than T_{max} , the temperature where $G(T)$ has a maximum.

4.1. DEM spline method using Hinode EIS lines

For a few sample cases based on EIS lines we have used the method described in Del Zanna (1999), where the *DEM* is assumed to be a spline function. The *DEM* is well constrained up to $\log T[K]=6.5$ (3 MK) with strong iron lines, up to Fe xvi. EIS observes several weak Fe xvii lines that can be used to constrain the higher temperatures. Note that all the Fe xvii lines are blended to some degree (see Del Zanna 2008 and Del Zanna & Ishikawa 2009 for details on line identifications and atomic data). For example, the strongest Fe xvii line at 204.66 Å with a cool line, identified as partly due to Fe xviii at 204.704 Å in Del Zanna (2009b). It turns out that the Fe xvii lines, despite having a $\log T_{\text{max}} = 6.75$, have $\log T_{\text{eff}}=6.5$ in the cores of active regions, given the large *DEM* at 3 MK. So, effectively, the Fe xvii lines provide an upper limit to *DEM* values above 3 MK.

A further upper limit around 10 MK is obtained from the Fe xxiii 263.8 Å line (see Del Zanna 2008; Del Zanna et al. 2011b for a discussion on blending for this line).

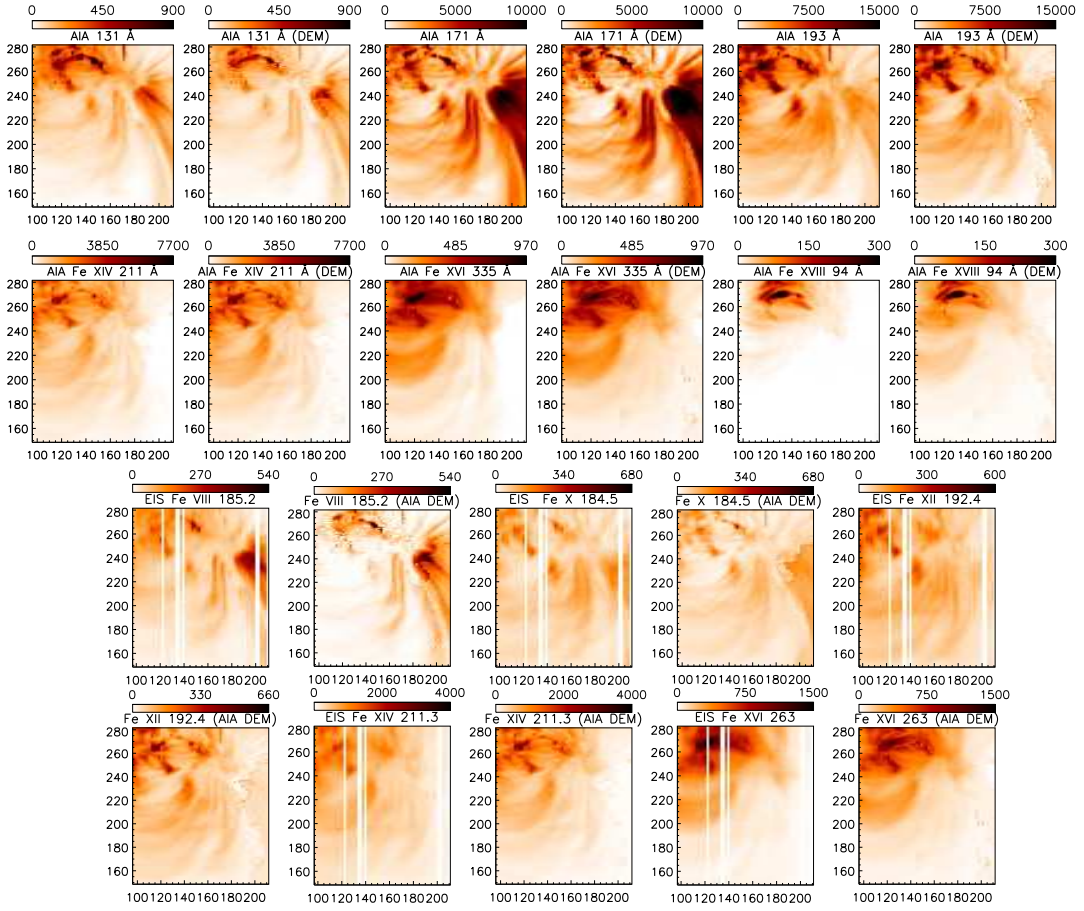


Fig. 4. Top two rows: observed and estimated (from the AIA *DEM* modelling) AIA images (negative, DN/s) of the on-disk AR observation of 2010 Oct 26, rebinned to the EIS resolution. Lower two rows: monochromatic images (negative, photons $\text{cm}^{-2} \text{s}^{-1} \text{arcsec}^{-2}$) in EIS lines, observed and estimated from the AIA *DEM* modelling.

A number of calcium lines are often used for emission measure purposes (Warren et al. 2012), however, a number of problems have been found here. Most of the Ca xiv and Ca xvi lines were found blended. The strongest Ca xv 200.97 Å line (Del Zanna 2012b), formed around 4.5 MK, has intensities always at odds with those of the other lines. The Ca xvii 192.82 Å is the strongest line formed around 5–6 MK observed by EIS (Del Zanna 2008), but unfortunately is severely blended. The main blending is due to the strong Fe xi 192.813 Å line. Its intensity is estimated here via a branching ratio (0.21 photons) with the 188.216 Å line (Del Zanna et al. 2010b). Several strong O v lines are also blending Ca xvii (Young et al. 2007). Following our benchmark work (Del Zanna 2009a), the O v 192.9 Å emission was used to estimate the contribution to the 192.8 Å blend. We often found no measurable Ca xvii emission, in which case an estimate of the Ca xvii minimum intensity was considered in the *DEM* modelling as an upper limit.

EIS also observes several lines from Ni, Ar, S, Si that can be used to measure quite accurately their coronal abundances (relative to iron), as described in Del Zanna (2012b) and discussed below. This is an important feature of the EIS spectra. A few weak lines from K, P, Cr, Mn can also provide a rough estimate for these elements, as shown below.

4.2. Automatic *DEM* inversion

The *DEM* spline method requires significant user interaction and is not suitable to calculate *DEMs* on a pixel-by-pixel basis. We have tested various inversion methods on several AIA observations adopting a range of parameters. In particular, we have tested two among the most refined methods: the MCMC method of Kashyap & Drake (1998) and the regularised inversion of Hannah & Kontar (2012). Once the *DEM* was obtained, the predicted AIA count rates were then compared to the observed ones. We also predicted EIS radiances and compared them to simultaneous observations, whenever available. We found the overall results very unsatisfactory. The MCMC method failed in most places, although we notice that it works relatively well when emission lines are considered (Del Zanna et al. 2011c). A similar conclusion was obtained by Testa et al. (2012a), where the MCMC method was applied to recover the results of a 3-D simulation. The Hannah & Kontar (2012) code performed relatively well but in places failed completely to provide an acceptable representation of the observations.

Therefore, we had to develop a new method to obtain *DEMs* from the AIA EUV bands. As pointed out in Del Zanna et al. (2011c), part of the problem in using the AIA EUV data lies in their broad temperature response. Another reason why sometimes inversion methods fail is related to the multiple peaks that are present in the 94, 131, 211, and 335 Å bands. This problem of the multiple peaks has been partially solved with the above assumptions, i.e. by isolating the Fe xiv and Fe xvi in the 211, and

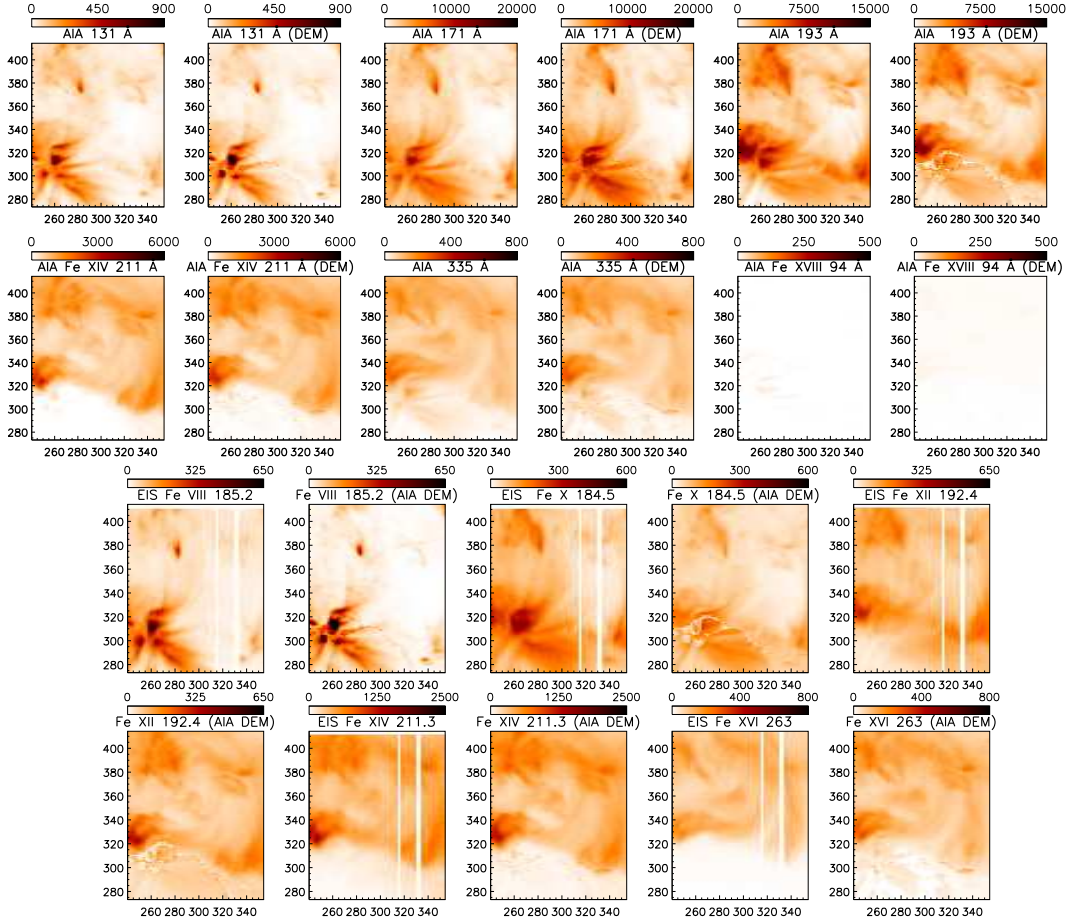


Fig. 5. Top two rows: observed and estimated (from the AIA *DEM* modelling) AIA images (negative, DN/s) of the on-disk AR observation of 2010 Nov 23, rebinned to the EIS resolution. Lower two rows: monochromatic images (negative, photons $\text{cm}^{-2} \text{s}^{-1} \text{arcsec}^{-2}$) in EIS lines, observed and estimated from the AIA *DEM* modelling.

335 Å bands, remove the hot component in the 131 Å band, and only use the 94 Å band in the hot regions where Fe XVIII emission can be isolated.

We chose a functional form for the *DEM* which allowed for the presence of near-isothermal structures, but was also multi-thermal and continuous. After various tests, we obtained the best results by representing the *DEM* in each AIA pixel as a superposition of nearly isothermal distributions, modeled as Gaussian functions of narrow width, and spaced across the entire temperature range. A similar approach was adopted by Aschwanden & Boerner (2011) to study single loops in one AR. Typically, we used about 10 Gaussians to cover the $\log T[\text{K}] = 5.6\text{--}6.6$ range, plus a few extra ones in the regions where hot emission is present. The widths of these Gaussians should be small to allow for the presence of near-isothermal plasma, but cannot be arbitrarily low, because of the limited accuracy by which the line emissivities are known (see also Judge 2010). These in turn mainly depend on the ionisation and recombination rates. After various tests, we found that a width of about 0.07 in $\log T[\text{K}]$ is a good compromise, resulting in a smooth *DEM* distribution.

For each pixel, a simple iterative procedure was applied by minimizing the χ^2 by varying the heights of the Gaussians. Values with very low intensities are typically not considered for the minimization. A variant of this method has been coded to be used for the Hinode EIS spectral lines.

Any inversion code is time consuming. For a small sample of 9600 points, MCMC required 5100s, the Hannah & Kontar

(2012) 2400s, and the present code a similar amount, although it can be reduced with a careful selection of parameters.

If the *DEM* is effectively a smooth distribution, a very good approximation for it is given by

$$DEM_L = \frac{I_{\text{ob}}}{Ab \int_T G(N, T) dT} \quad (7)$$

where the integral is performed over the entire temperature range where a line is formed. We use this approximation as a starting solution for the inversion, assuming for the AIA 131, 171, 193, 211, 335, 94 Å channels the temperatures (log values): 5.65, 5.85, 6.2, 6.3, 6.4, 6.8, and then interpolating them over a 0.01 temperature grid (in $\log T$). Note that such approximation, used in Widing & Feldman (1989) and following papers to estimate relative elemental abundances, fails completely when the plasma is nearly isothermal (Del Zanna et al. 2001; Del Zanna 2003b; Del Zanna et al. 2003).

An upper limit to the *DEM* distribution is easily obtained by considering the emission measure loci. The loci of the curves $I_{\text{ob}}/(Ab \times G(N, T))$ represent, at each temperature, the upper limit to the value of $EM_h(T, \Delta T)$. The limit is obtained only if the plasma is strictly isothermal at that temperature. The EM Loci method was applied for the first time by Strong (1978) and Veck et al. (1984) to the analysis of solar X-ray flare spectra, and later by Del Zanna & Mason (2003) to SOHO CDS observations

of solar active regions. The method has since become very common in the literature.

The upper boundary for the *EM* has been converted to an upper boundary for the *DEM* distribution by assuming a constant *DEM* in each temperature bin. As a lower boundary, the DEM_L values reduced by two orders of magnitude were chosen.

We have calculated the instrument responses $G(N, T)$ in the $T[\text{K}]=5-7.5$ range, but have limited the inversion to a more restricted temperature range, depending on the type of observations considered. The results of the inversion are quite sensitive to this choice. The *DEM* values in the $\log T[\text{K}]=5-5.5$ range are unconstrained by AIA, so we have adopted typical values in the range $10^{20.7--20.0} \text{ cm}^{-5} \text{ K}^{-1}$ for on-disk observations.

5. Validating the *DEM* method with AIA and EIS simultaneous observations

We have tested the *DEM* method on a few AIA and EIS simultaneous observations of active regions. We provide here two examples. Fig. 4 shows the result of the *DEM* modelling for the on-disk AR observation of 2010 Oct 26 obtained from the AIA data. The agreement between observed and estimated AIA count rates is truly remarkable. The *DEM* was also used to predict the Hinode EIS radiances (in photons $\text{cm}^{-2} \text{ s}^{-1} \text{ arcsec}^{-2}$), which are compared to the observed ones in the same figure. The agreement is surprisingly good, both in terms of absolute and relative values. The upper part was difficult to model, because hot emission from Fe xviii is present. The results for the on-disk AR observation of 2010 Nov 23 are shown in Fig. 5. The lower part was particularly difficult to fit because the hotter lines are very weak, while the cooler ones are very strong. This is a region where other methods we have tested typically failed.

In both cases, we adopted the coronal abundances that are discussed below. The emissivities were calculated with the present atomic data (listed in the beginning of Sect. 3) and a constant electron density of $2 \times 10^9 \text{ cm}^{-3}$, a typical average AR value. The excellent agreement confirms the validity of the AIA and EIS calibrations, as well as the accuracy of the atomic data. The comparisons also indicate a good reliability of the procedure, which provides confidence in the results that are obtained with the full-resolution AIA data, described below.

Obviously, the fact that in most pixels the method is able to reproduce AIA count rates and EIS radiances within about 30% does not necessarily mean that the obtained *DEM* is an exact solution. Indeed, as mentioned, the inversion method is ill-posed and there is no unique solution. Also, broad-band imaging can never replace the accuracy and amount of information that can be obtained with spectroscopic measurements.

6. The main contributions to the AIA EUV bands in the cores of ARs

To assess the various contributions to the AIA bands, a large region within the core of the AR observed on the 2010 Oct 26 was selected (see Fig. 4), to avoid possible small misalignments between EIS and AIA, and increase the S/N in the EIS spectra. Table 1 shows the observed AIA count rates, together with the predicted ones for the 193 and 211 Å bands, obtained by convolving the calibrated EIS spectra with the AIA (ground calibration) effective areas. Good agreement is found. A large uncertainty in the AIA DN estimated directly from EIS is related to the measurement of the EIS bias, done here by taking average values around 205.5, and 210.5 Å in regions free of strong lines,

for the 193 and the 211 Å bands. Good agreement is found. The EIS 171 Å line is very uncertain (50%) because it is at the limit of the detector and has very low counts.

Table 1. Observed and predicted (from the EIS data and the *DEM* modelling) SDO/AIA count rates (averaged DN/s per EIS pixel) for the 2010 Oct 26 (on-disk) region.

Band (Å)	Obs. (AIA)	Pred. (EIS)	Pred. (DEM)
94	45	-	45(0%)
131	169	-	127(-25%)
171	2385	-	2790(17%)
193	4913	5570(13%)	4694(-4%)
211	2610	2382(-9%)	3385(30%)
335	339	-	368(10%)

Fig. 6 (top) shows the *DEM* obtained with the spline method and the elemental abundances discussed below. We used this *DEM* to calculate line and continuum emission for all the AIA bands and calculate the total AIA simulated count rates, displayed in Table 1. Very good agreement is found.

In the 171 Å band, the main Fe ix line contributes 1930 DN/s, while about 10% (310 DN/s) comes from Ni xiv 171.37 Å. In the 193 Å band, about half comes from the Fe xii lines. In the 211 Å band, the Fe xiv 211.3 Å line contributes about 50% of the predicted counts. In the 335 Å band, the Fe xvi 335.4 Å line contributes about 80%.

A significant fraction of the counts in the 94 and 131 Å bands comes from the free-bound continuum, formed around 3 MK. With the present abundances, the continuum contributes about 10% and 36%, respectively. With photospheric abundances, the contributions increase to about 20% and 50% for the 94 and 131 Å bands.

6.1. The AIA 94 Å band and Fe xviii

The AIA simulated count rates as a function of temperature and wavelength are shown in Fig. 6. A significant contribution comes from Fe xviii. After Fe xviii, the newly-identified Fe xiv 93.61 Å line produces the main contribution to the band, while minor contributions come from Fe x, and the set of lines from Fe ix and Fe xii discussed previously.

It is interesting to note that the Fe xviii intensity is due to the 3 MK plasma, as Fig. 6 (middle) shows. This is caused by the large *DEM* peak at 3 MK, a temperature where the Fe xviii contribution function is only about 8% its peak value at 7 MK. This is a surprising result, although it somewhat uncertain because it relies on the accuracy in the ion abundance toward the wings. Note that the Ca xvii 192.8 Å line has instead at 3 MK a predicted emissivity of only 1% its peak value at 6 MK. This is consistent with the fact that we found several cases where this line is not observed when instead Fe xviii is expected to be present. This is another somewhat surprising result, considering that Ca xvii has a lower peak formation temperature (6 MK) than Fe xviii (7 MK).

Obtaining the Fe xviii emission with *DEM* modelling based on EIS lines is clearly difficult, especially when Fe xvii lines are not measurable. A better way to estimate the Fe xviii contribution to the AIA 94 Å band is to subtract all the known contributions

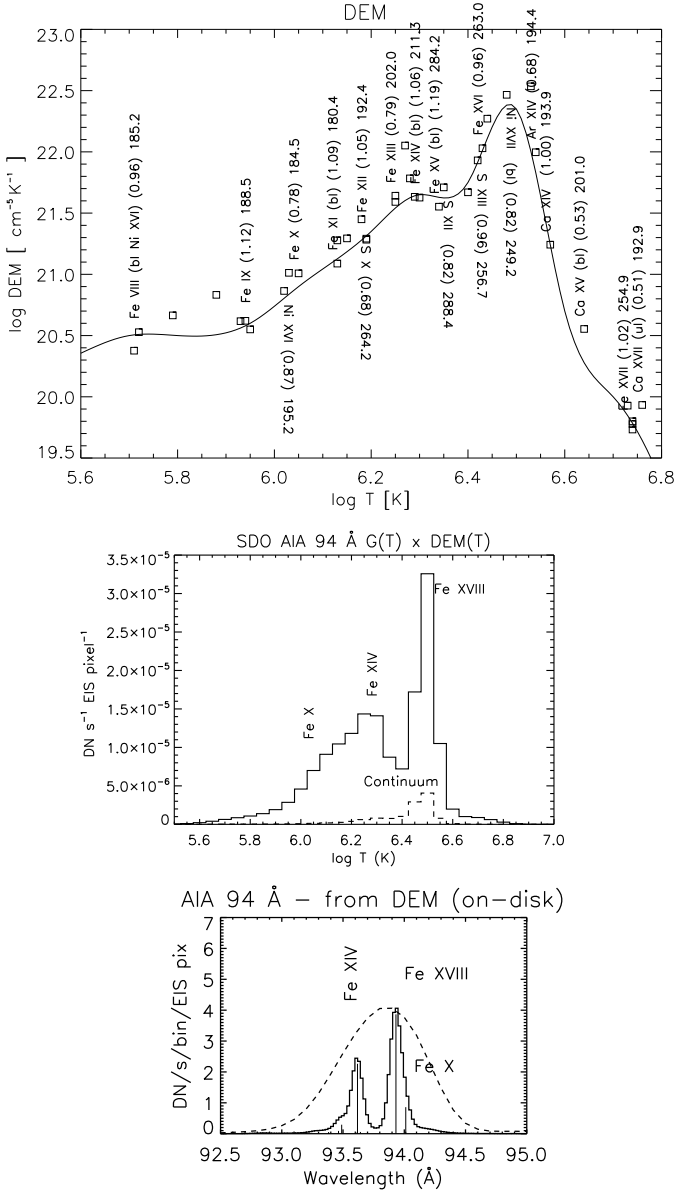


Fig. 6. Above: The *DEM* of the 2010 Oct 26 active region core. The numbers in parentheses are the ratios between predicted and observed intensities. The points are plotted at the temperature T_{max} , and at the theoretical vs. the observed intensity ratio multiplied by the *DEM* value. Below: corresponding AIA 94 \AA count rates as a function of temperature and wavelength (the dashed curve is the normalised AIA effective area)

from the observed count rates. The Fe VIII–Fe XIV contributions can in fact be estimated quite accurately using the EIS observations. We applied our automatic *DEM* inversion for each EIS pixel using 12 iron lines. We then calculated the contributions of the various ions and the continuum for each EIS pixel, in AIA 94 \AA DN/s. The results for the two selected ARs are plotted in Fig. 7. Panels 7 in the same figure show the observed count rates with all the various known contributions subtracted. The residual emission is expected to be due mostly to Fe XVIII, although at least one relatively strong emission line is still unidentified. The lower panel (Nov 23 observation) indicates that the unidentified line has a low formation temperature, close to Fe VIII.

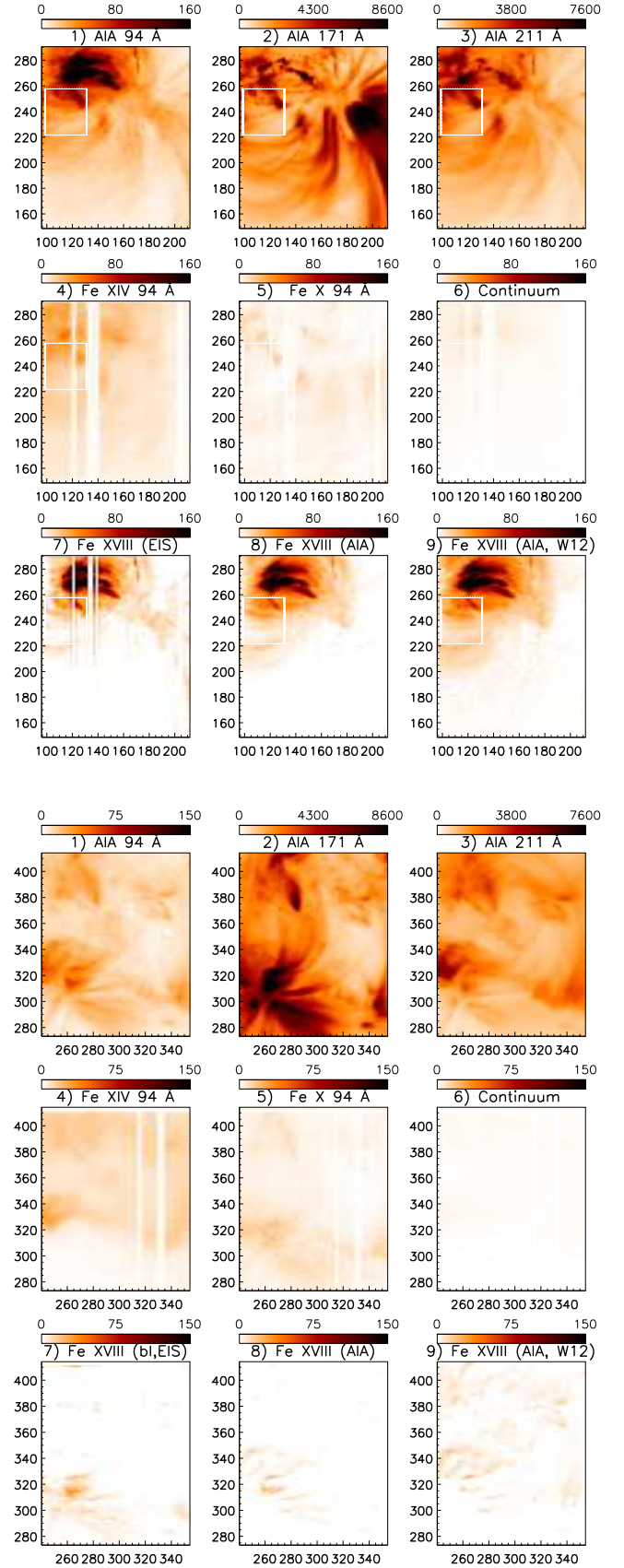


Fig. 7. AIA images (negative, DN/s/pixel) of the on-disk AR observations of 2010 Oct 26 (above) and Nov 23 (below), rebinned onto the EIS resolution. Panels 4,5,6 show the 94 \AA predicted count rates in Fe XIV, Fe X, and the continuum. Panels 7 and 8 show the residual count rates in the AIA 94 \AA (due to Fe XVIII), estimated from EIS and AIA. Panels 9 show the Fe XVIII estimate of Warren et al. (2012).

It is clear that the newly-identified Fe xiv 93.61 Å always provides an important contribution to the band, the dominant one whenever Fe xviii emission is not present. A significant contribution also comes from the continuum, while a minor one from Fe x 94.012 Å. The Fe viii, Fe ix, Fe xii lines produce a negligible contribution. This is a new and important result. There is good agreement between the present and the Warren et al. (2012) method to estimate the Fe xviii emission, as shown in Fig. 7.

7. The multi-thermal emission at the AIA resolution

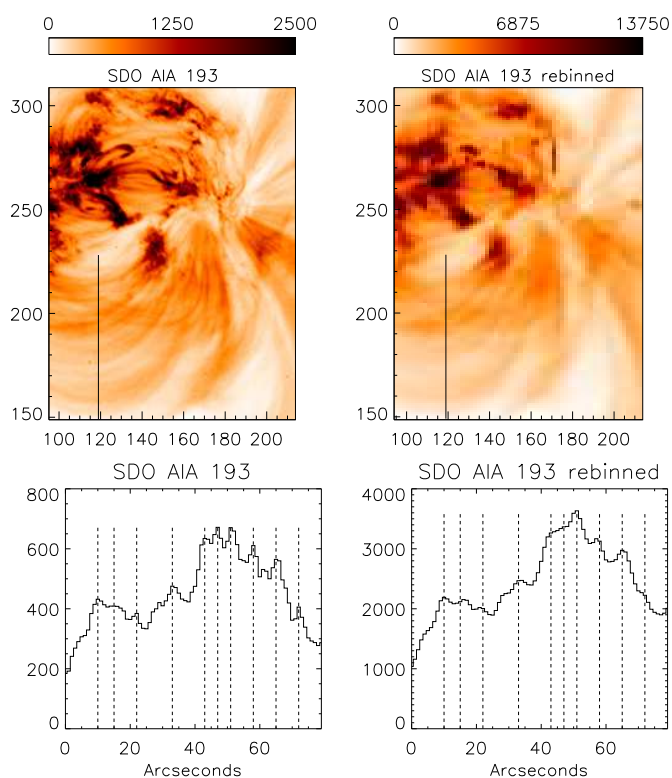


Fig. 8. Above: a zoom up of a full-resolution AIA 193 Å image (left, negative, at 11:40 UT) of the 2010 Oct 26 active region within the EIS FOV, and the corresponding rebinned image obtained by converting the AIA observations into the EIS spatial and temporal resolution (right, negative). Below: profiles of the count rates along the N-S segment shown in the images. The locations of a few loop structures that appear as nearly resolved in the full-resolution AIA 193 Å image are indicated with dashed lines.

One of the aims of this paper was to discuss the multi-thermal emission in the AR cores at the higher AIA resolution. To illustrate how significant the increase is, we consider the 2010 Oct 26 AR observations. Hinode EIS observed the AR during 10:50–11:50 UT. We compare in Fig. 8 a full-resolution AIA 193 Å image taken at 11:40 UT with the rebinned image, obtained by converting the AIA observations into the EIS spatial and temporal resolution (the same already shown previously). We corrected for the AIA stray light using the results of Poduval et al. (2013). We note that the correction for the stray light is significant (up to 50% changes) for some active regions and enhances the contrast between the loop structures and the background.

Fig. 8 also shows profiles of the count rates along a segment shown in the images. Several loop structures appear nearly re-

solved by AIA (they are indicated with dashed lines), but many of them become hardly visible at the EIS resolution. We note that there is a region free of loop structures, near the AR core (at pixel position 80), where the AIA count rates are similar to those far away from the AR (at pixel position 0, the bottom of the image). In the rebinned image, the count rates in this core region are instead almost twice the values at the bottom of the image. This clearly illustrates that a factor of two increase in the ‘unresolved’ background observed by Hinode EIS (and SOHO CDS) in AR cores can simply be caused by its lower resolution.

We then performed the automatic AIA *DEM* modelling on the full-resolution AIA data of the active region using the set of abundances discussed below. To increase the S/N, we averaged the AIA images during a minute, between 11:40 and 11:41 UT. We corrected for the AIA stray light using the results of Poduval et al. (2013) and applied our corrections for the AIA plate scale. We then calculated the predicted radiances in a sample of EIS lines. Sample results are shown in Fig. 9, centered on a bundle of 3 MK loops, clearly seen in AIA 335 Å (Fe xvi). These images, together with the intensity profiles across the loops, clearly show the presence of a multitude of loops that are mostly not co-spatial at the various temperatures, and become more unresolved at higher temperatures. The locations of a few cool loops which appear as nearly resolved in Fe viii and Fe ix are shown in the profiles of Fig. 9 as dashed lines. Higher-temperatures loops in e.g. Fe xii are normally not co-spatial. In each single AIA pixel, there is clearly some superposition of different structures along the line of sight.

The ‘contrast’ of the structures as seen in single emission lines is clearly enhanced, compared to the original AIA images, which confirms that the multi-thermal nature of the AIA bands is somewhat ‘blurring’ our view of the AR structures. For example, we notice the significant differences between the AIA 131 Å image and the reconstructed Fe viii towards the AR core (upper images, 100–150'' in the profiles). The *DEM* modelling predicts very little unresolved background emission in the AR core in lines formed up to 1 MK (as Fe x). The unresolved background emission is instead about a factor of 3–4 higher in Fe xii–Fe xiii, and much higher in Fe xiv–Fe xvi, compared to the values outside the core (in the bottom of the image). This confirms the TRACE and SOHO CDS results of Del Zanna & Mason (2003), although the higher AIA resolution now shows much lower levels of unresolved emission, relative to the total emission (less than 50% instead of 80%). The increased spatial resolution clearly allows us to resolve more and more structures and enables us to better ‘see through’. One question naturally arises: what is this unresolved emission? It is likely composed of a multitude of loops of sizes well below the resolution of current instrumentation. We believe that a factor of 2–3 increase is just due to an increase in the iron abundance, rather than a density effect, as discussed below.

There is an unresolved 3 MK bundle of loops seen in Fe xvi, with footpoints approximately aligned along the E-W direction. The bundle of 3 MK loops is also clearly visible with EIS (cf. Fig. 4). We note that these loops are also clearly visible in Fe xiv, although most of the emission in this ion is formed at 3 MK and not at its peak formation temperature of 2 MK, as discussed below.

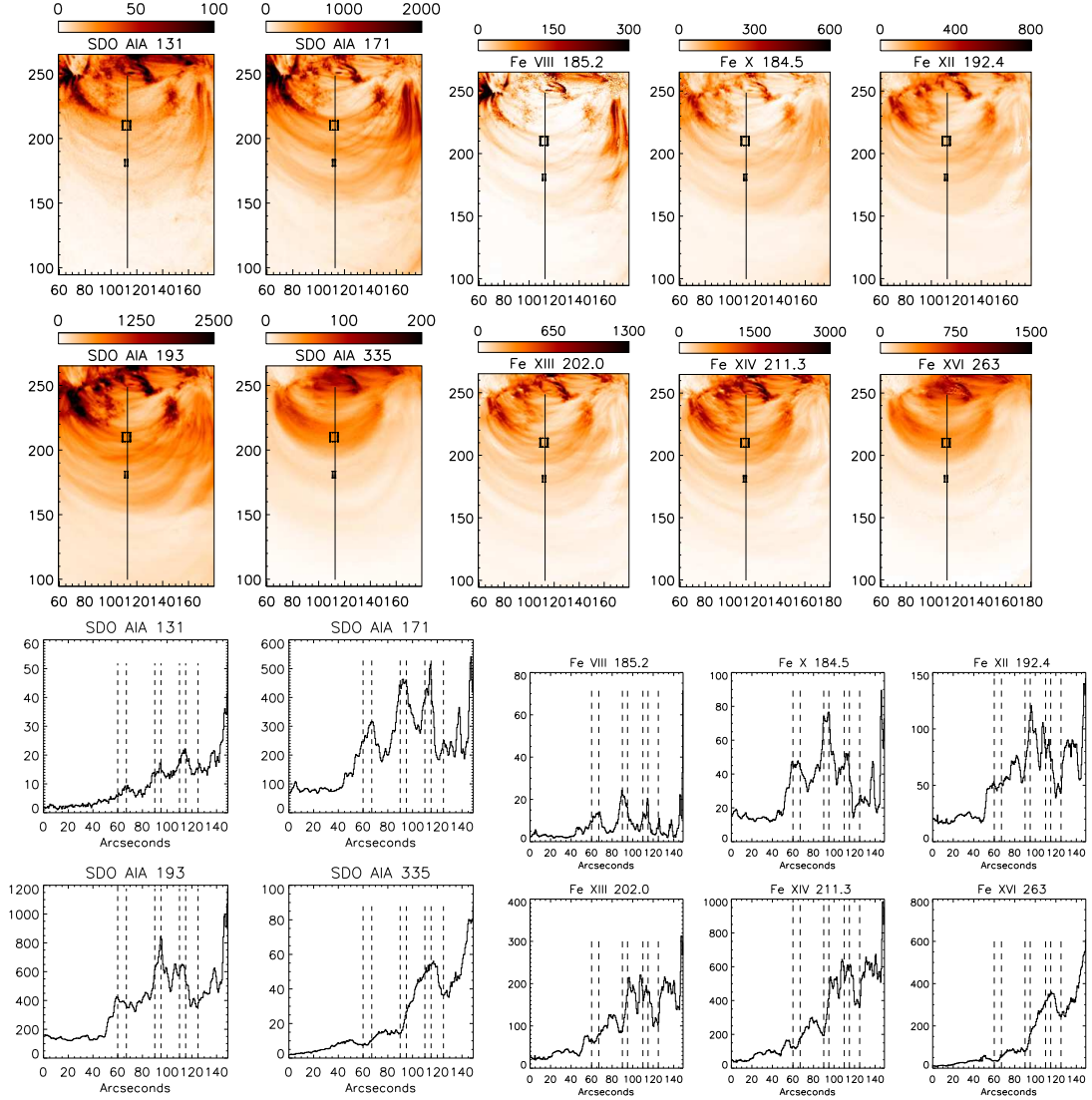


Fig. 9. Top left: The full-resolution AIA images of the 2010 Oct 26 active region. The two selected regions are shown (the larger on the 3 MK loops, the smaller for the background), as well as a cut across the 3 MK loops. Top right: the corresponding radiances in a few EIS lines, predicted from the AIA *DEM* modelling. Below: profiles of the count rates and radiances across the 3 MK loops.

8. Elemental abundances and path lengths for the 3 MK loops

A region within the centre of this 3 MK bundle of loops (at coordinates 110, 210, see Fig. 9) with no underlying obvious emission from e.g. moss regions was chosen for detail analysis using the Hinode EIS spectral lines. We obtained averaged EIS calibrated spectra, and measured the average electron density from the Fe XIV lines and found $10^{9.2} \text{ cm}^{-3}$. We note that the Fe XIV lines are not formed at peak ion abundance, so the densities need to be obtained by calculating the theoretical ratios at 3 MK, something that is not normally considered in the literature. We obtained consistent results from other ions.

We performed a *DEM* modelling using the EIS iron lines, adopting at first the photospheric abundances of Asplund et al. (2009). We then calculated the path length as described in Section 4, by integrating the *DEM* in the $\log T[\text{K}] = 6.3\text{--}6.6$ range and considering a density of $10^{9.2} \text{ cm}^{-3}$. We obtained $350''$, an unreasonable number, considering that the size of the bundle of loops as seen in the plane of the sky is only of the order of $20''$ (see e.g. Fig. 9). Considering that the EIS calibration

is now quite well established (Del Zanna 2013), and the atomic data for the iron lines are quite accurate now, there are only two improvements that can be applied to reduce the path length. The first one regards ‘background/foreground’ subtraction. Clearly, the 3 MK bundle of loops is embedded in an environment rich of loops at lower temperatures, plus some line-of-sight emission at 3 MK. Subtracting this emission lowers the observed emission measure. Given the EIS resolution and FOV, there is a limited choice, so a small nearby region (where Fe XVI emission is low) was chosen (at coordinates 110, 180, see Fig. 9). Subtracting the emission from this region not only reduces the emission measure, but also increases the averaged density (which also reduces the path length), to a value of about $10^{9.5} \text{ cm}^{-3}$, providing a path length of $54''$.

The electron density was obtained using the ‘emissivity ratio’ technique (see Del Zanna et al. 2004 for details), which divides the observed intensity of a line (I_{ob}) by its emissivity:

$$F_{ji} = \frac{I_{\text{ob}} N_e C}{N_j(N_e, T_e) A_{ji}} \quad (8)$$

Table 2. The present hot core loops abundances

El.	Phot.	Coronal	Present	FIP (eV)
H	12	12	12	13.56
C	8.43	8.59	-	11.26
N	7.83	8.0	-	14.53
O	8.69	8.89	-	13.61
Ne	7.93*	8.08	-	21.56
Mg	7.6	8.15	8.10	7.64
Al	6.45	7.04	6.95	5.98
Si	7.51	8.10	8.00	8.15
P	5.41	-	5.41	10.49
S	7.12	7.27	7.17	10.36
Ar	6.40*	6.58	6.30	15.76
K	5.03	-	5.53	4.34
Ca	6.34	6.93	6.87	6.11
Cr	5.64	-	6.14	6.78
Mn	5.43	-	5.93	7.43
Fe	7.50	8.10	8.00	7.87
Ni	6.22	6.84	6.53	7.63

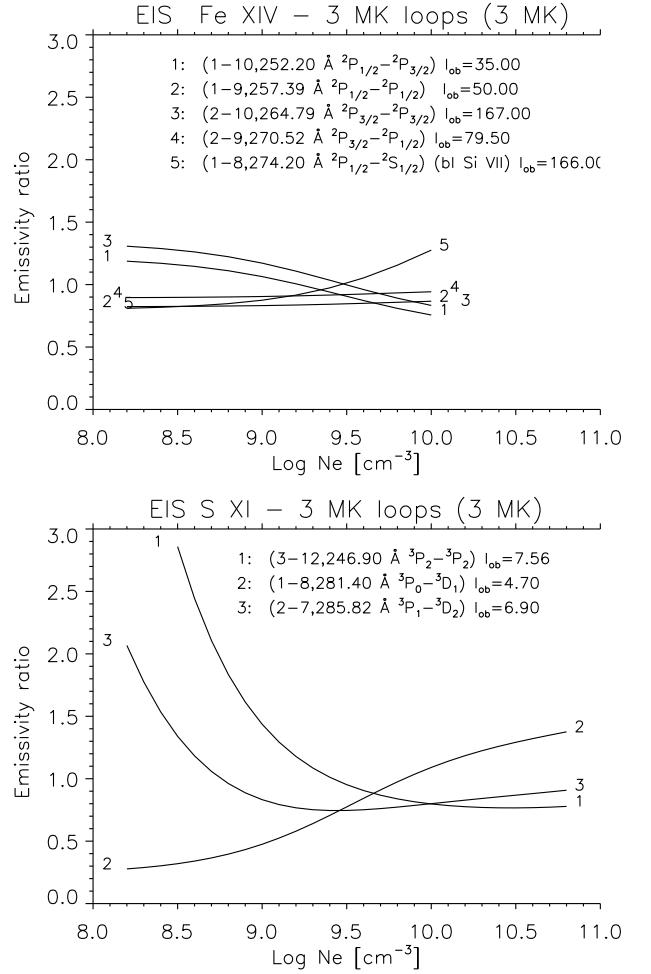
Notes. Photospheric (Phot) abundances are from Asplund et al. (2009), the coronal ones from Feldman et al. (1992). Ar and Ne do not have photospheric measurements.

calculated at a fixed electron temperature T_e (at 3 MK) and plotted as a function of the density N_e (C is a scaling constant). N_j is the population of the upper level, while A_{ji} is the transition probability of spontaneous emission. Only Fe xiv and S xi provide useful diagnostics because the lines are strong in a 3 MK plasma. The two ions provide similar densities, as shown in Fig. 10. There are other density-sensitive Ni xv and Ca xv lines but they are weak, confirming that the density is not higher than $10^{9.5} \text{ cm}^{-3}$.

The reduction in the path length is significant but not enough, so reasonable path lengths are obtained only by increasing the iron abundance. We adopted a reasonable increase of 0.5 dex (a factor of 3.16 increase), providing a path length of $17''$. This is a reasonable value, close to the overall size of the bundle of loops, and that implies a spectroscopic filling factor close to one. We note that the factor of 3.16 increase is consistent with typical X-ray measurements of the relative O/Fe and Ne/Fe ratios (which will be discussed in a separate paper), assuming that the high-FIP elements have the photospheric (Asplund et al. 2009) abundances.

Having fixed the iron abundance to 8.0 dex (on the usual logarithmic scale where hydrogen is 12), we then used the DEM modelling to find the abundances shown in Table 2. Fig. 11 shows the DEM obtained with the spline method from the background-subtracted intensities. For comparisons, it also shows the DEM obtained from the same set of abundances but without subtracting the background. The DEM is very well constrained by strong lines from all the iron ionisation stages from Fe viii to Fe xvi. The Fe xvii lines provide a constrain at higher temperatures. The DEM is nearly isothermal, centred at 3 MK, and with a peak value two orders of magnitude higher than the emission measures at lower or higher temperatures.

Table 3 summarises the results for the strongest lines. Excellent agreement between theory and observations is found, although we remind that discrepancies of factors of two are present if the EIS ground calibration is adopted. The abundances of K, Ar, Ca, P, S, and Ni (relative to iron) are well constrained, since the K xiii, Ar xiii, Ar xiv, Ca xiv, P xii, S xii, S xiii, Ni xv,


Fig. 10. Emissivity ratio curves relative to the main Fe xiv and S xi transitions observed by Hinode EIS on the bundle of 3 MK loops of Fig. 9.

Ni xvi, and Ni xvii lines are all formed around 3 MK (see the effective temperatures in Table 3). We recommend that the following lines are used for future studies: K xiii 208.109 Å, Ar xiv 194.396 Å, Ca xiv 193.866 Å, P xii 278.286 Å, S xii 288.421 Å, S xiii 256.685 Å, Ni xv 176.710 Å, Ni xvi 195.271 Å, Ni xvii 249.185 Å. The P xii 278.286 Å is weak but is the resonance line for this ion, so effectively the EIS measurement provides an upper limit to the P abundance. A similar argument applies to the K xiii 208.109 Å resonance line and the Cr ix 210.611 Å line. The strong Ca xv 200.971 Å is at odds with all the other ones, indicating possible problems in the calcium ion abundances or with the atomic data for this ion, or the presence of an unknown blend.

The lower-temperature Ar xi is at odds with the other ions. This could partly be due to different abundances in the lower-temperature plasma, as the S x abundance seem to suggest as well. The present abundances differ significantly from previous ‘canonical’ coronal and photospheric abundances, as shown in Table 2. We note that we obtained a sulfur abundance close to the photospheric value suggested by Asplund et al. (2009), which means that the 3 MK loops had a FIP enhancement of 3.16. Sulphur has in fact a relatively high FIP and its abundance variations follow those of the higher-FIP elements. The abundances of Ar and P (two high-FIP elements) are also close to the values suggested by Asplund et al. (2009), although we note that

Table 3. Observed and predicted radiances for the 3 MK loops.

λ_{obs}	I_{obs}	T_{max}	T_{eff}	R	Ion	λ_{exp}	r
176.69	18.9	6.40	6.47	1.06	Ni xv	176.710	0.99
177.23	40.8	6.05	6.11	0.97	Fe x	177.240	1.00
180.39	62.9	6.13	6.17	1.06	Fe xi	180.401	0.99
183.45	9.1	6.56	6.50	0.94	Ar xiv	183.450	0.13
					Ca xiv	183.460	0.87
184.53	19.9	6.05	6.13	0.88	Fe x	184.537	0.94
185.22	40.7	5.73	6.39	0.89	Ni xvi	185.230	0.74
					Fe viii	185.213	0.25
185.66	2.1	6.40	6.46	1.10	Ni xv	185.696	0.88
					Ni xv	185.690	0.06
186.61	27.4	5.72	6.39	0.76	Ca xiv	186.610	0.70
					Fe viii	186.598	0.29
188.21	40.9	6.13	6.17	0.86	Fe xi	188.216	1.00
188.49	8.4	5.95	6.06	0.88	Fe xii	188.450	0.08
					Mn ix	188.480	0.06
					Fe ix	188.497	0.83
188.80	6.0	6.27	6.44	0.43	Ar xi	188.806	0.97
189.93	3.7	5.94	6.06	1.07	Fe xii	189.925	0.09
					Fe ix	189.941	0.89
191.40	4.5	6.51	6.48	0.79	Ar xiv	191.404	0.62
					Fe xv	191.412	0.34
192.39	21.4	6.19	6.25	0.96	Fe xii	192.394	1.00
192.81	9.1	6.13	6.16	0.88	Fe xi	192.813	0.95
193.87	23.2	6.57	6.50	0.96	Ca xiv	193.866	1.00
194.05	3.2	6.19	6.46	1.18	Ni xvi	194.046	0.88
					Fe xii	194.078	0.08
194.41	5.5	6.54	6.49	0.99	Ar xiv	194.396	0.99
195.11	65.4	6.19	6.25	1.00	Fe xii	195.119	0.99
195.26	7.8	6.02	6.45	0.95	Ni xvi	195.271	0.88
195.54	1.6	6.08	6.46	0.83	Ni xv	195.521	0.93
200.99	14.7	6.64	6.50	0.56	Ca xv	200.971	0.98
202.04	37.9	6.25	6.36	1.09	Fe xiii	202.044	1.00
204.65	5.2	6.74	6.50	1.05	Fe xvii	204.668	0.99
205.92	0.936	5.46	6.45	0.87	Ar xiii	205.950	0.81
					Fe xiii	205.914	0.07
208.12	4.0	5.50	6.48	0.51	K xiii	208.109	0.90
					Ni xv	208.154	0.07
208.31	4.6	6.19	6.43	0.82	Ca xv	208.321	0.65
					Fe xii	208.316	0.33
208.63	4.9	6.24	6.43	0.82	Ca xvi	208.604	0.51
					Fe xiii	208.667	0.39
					Cr viii	208.633	0.05
211.31	213.0	6.30	6.44	1.05	Fe xiv	211.317	1.00

Notes. λ_{obs} (Å) are the observed wavelengths, I_{obs} the background-subtracted Hinode EIS radiances (photons $\text{cm}^{-2} \text{s}^{-1} \text{arcsec}^{-2}$) obtained with our new EIS calibration (Del Zanna 2013), T_{max} and T_{eff} the maximum and effective temperature (log values, in K; see text), R the ratio between the predicted and observed radiances, Ion and λ_{exp} (Å) the main contributing ion, and r the fractional contribution to the blend.

photospheric measurements of argon do not exist, and remind that the P xii line is weak, so the P abundance is uncertain.

If we return now to the issue of the iron absolute abundance, we note that it could be higher than 8.0 dex (and the filling factor smaller than one), but then the abundances of the high-FIP elements would also have to be higher than photospheric. We also note that direct line-to-continuum measurements obtained from ground-based eclipse observations by Mason (1975) also suggest an increased calcium and iron abundance by a factor of about three, close to active regions. Mason (1975) measured the calcium abundance as 6.8 dex (i.e. only 15% lower than the value

Table 3. Contd.

λ_{obs}	I_{obs}	T_{max}	T_{eff}	R	Ion	λ_{exp}	r
246.20	19.1	6.25	6.35	1.02	Fe xiii	246.209	0.99
246.90	7.6	6.27	6.42	0.72	S xi	246.895	1.00
249.17	199.0	6.48	6.49	0.76	Ni xvii	249.185	1.00
251.06	137.0	6.43	6.49	0.86	Fe xvi	251.063	1.00
251.94	51.3	6.25	6.35	0.73	Fe xiii	251.952	0.99
252.19	35.4	6.29	6.44	1.02	Fe xiv	252.199	1.00
254.87	5.7	6.74	6.49	1.09	Fe xvii	254.885	0.94
					Fe xiii	254.905	0.06
256.67	200.0	6.42	6.48	1.04	S xiii	256.685	1.00
257.38	50.4	6.29	6.44	1.11	Fe xiv	257.394	0.99
261.05	9.2	6.15	6.25	1.20	Si x	261.056	1.00
262.98	225.0	6.43	6.49	0.93	Fe xvi	262.976	1.00
264.23	8.7	6.18	6.31	0.78	S x	264.230	0.98
264.78	167.0	6.29	6.43	0.96	Fe xiv	264.788	0.98
265.00	21.9	6.43	6.49	0.93	Fe xvi	265.000	1.00
270.51	79.5	6.29	6.44	1.01	Fe xiv	270.520	1.00
271.98	12.0	6.15	6.25	0.83	Si x	271.992	0.99
274.20	168.0	6.29	6.44	0.92	Fe xiv	274.203	0.99
275.36	8.4	5.79	5.92	0.63	Si viii	275.361	1.00
277.25	7.2	6.15	6.25	1.15	Si x	277.264	1.00
278.26	2.4	6.32	6.46	0.71	P xii	278.286	0.99
281.40	4.7	6.28	6.42	0.88	S xi	281.402	0.99
284.15	2510.0	6.34	6.47	1.05	Fe xv	284.163	1.00
285.82	6.9	6.27	6.41	0.94	S xi	285.822	0.98
288.15	14.4	6.44	6.48	1.02	Ni xvi	288.165	0.98
288.40	31.9	6.35	6.46	0.92	S xii	288.421	1.00

suggested here), and the iron abundance as 7.85 dex (only 30% lower than the present value).

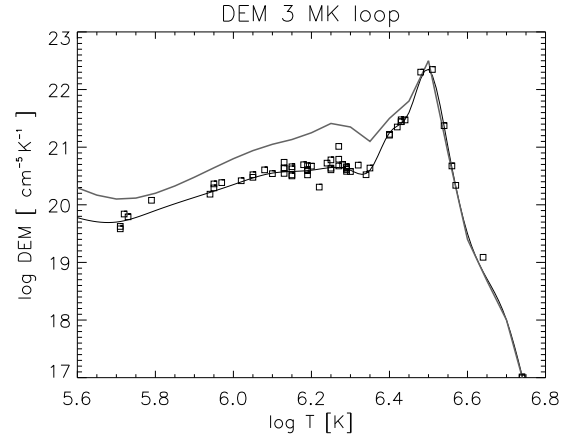


Fig. 11. The DEM of the 2010 Oct 26 active region 3 MK loops, obtained from the background-subtracted intensities, with the present coronal abundances. The points are plotted at the temperature T_{max} , and at the theoretical vs. the observed intensity ratio multiplied by the DEM value. The grey thick line shows the DEM obtained directly from the EIS intensities without background subtraction.

9. Hi-C

The most direct way to address the issue of the existence (or not) of sub-resolution structures in active regions is to actually take images and spectra at higher resolution. Hi-C, during its successful flight (Cirtain et al. 2013), obtained coronal images of

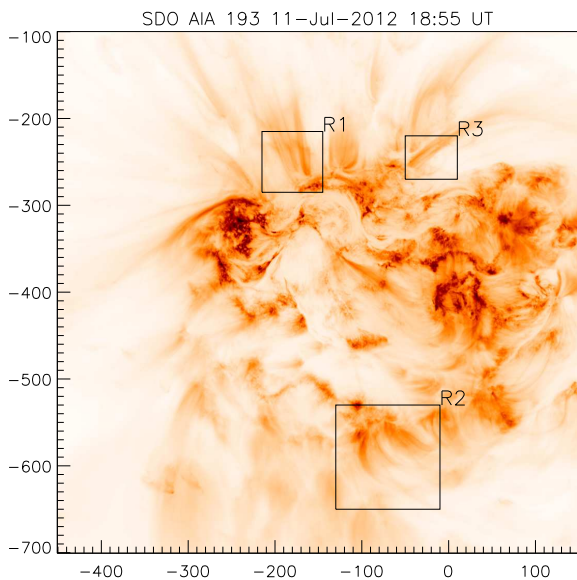


Fig. 12. The AR as observed in the AIA 193 Å band during the Hi-C rocket flight, with a few selected regions at the legs of several coronal loops.

an active region at the highest resolution to date (0.25'', about five times better than AIA) in a broad-band very similar to the AIA 193 Å one. This band does not show the hot core loops, but is useful to discuss the 1.5–2.5 MK background emission, still unresolved at the AIA resolution. Hi-C could in principle also provide information on coronal loops that emit at these temperatures, but unfortunately the FOV only included the footpoints and legs of a few warm loops, and not their full extent. The AR observed by Hi-C was very dispersed, had a large number of filaments, some very hot (10 MK) loops, and a large number of moss regions. The features discussed in Certain et al. (2013) were probably low-lying ones associated with a large filament, and not clearly with coronal loops.

A direct comparison of the Hi-C and AIA 193 Å images immediately shows the much higher resolution of Hi-C, but only in the low-lying (e.g. not properly ‘coronal’) moss regions, and in the low-temperature filament structures, since the prominence-corona transition region emits strongly in cool ($T < 1$ MK) lines observed by the 193 Å band. That filaments have a very fine structure has been known from ground-based observations. The fine structure of the moss emission confirms the fine-structure of the loops that are connected to it.

In all the other regions, very little differences between the Hi-C and AIA 193 Å images are present, indicating that any unresolved emission in the AIA images is still unresolved at the Hi-C resolution. Some examples are given below.

First, we have processed the AIA full-Sun images, adjusting their plate scale and correcting for the stray light, as described previously. We then averaged the AIA images during a minute, at 18:55 UT. We took the simultaneous Hi-C level 1.5 data (as available from the Virtual Solar Observatory), and averaged the exposures obtained at 18:55:02, 18:55:08, and 18:55:13 UT. We applied the suggested rotation (-1.81969 degrees, keyword `roll_angle`), and then checked the co-alignment of the Hi-C data with the AIA 193 Å image by plotting one image over the other one in a few sample moss areas. The moss areas are the best because are clear low-lying bright features. We found the need to apply a small shift of 1'' towards the W direction and 1.5'' towards the N direction.

Considering the multi-thermal nature of the 193 Å band discussed previously, a proper physical interpretation of the Hi-C 193 Å images requires at least some *DEM* modelling. We therefore performed the *DEM* inversion on the full Hi-C FOV using the AIA data, adopting the present coronal abundances and a constant electron density of $2 \times 10^9 \text{ cm}^{-3}$, a typical average AR value.

We then selected three regions (R1, R2, and R3 shown in Fig. 12) where some coronal loop legs are present. Fig. 13 shows region R1, which contains several loop legs that appear resolved in the 131 Å (Fe VIII) and 171 Å (Fe IX) images. The positions of seven cool loops are indicated with dashed lines in the profiles in Fig. 13. According to the *DEM* modelling, these cool loops are emitting in Fe VIII and Fe IX, and are still visible around 1 MK (see the Fe X image), but at higher temperatures (see the Fe XII image) they disappear. Instead, other structures appear in the line of sight at temperatures at and above 1 MK. The strongest cool emission is concentrated between 10 and 25'', where the *DEM* modelling predicts a depletion of the hotter Fe XII and especially Fe XIV. Such thermal structuring in active regions is common (although not ubiquitous). It was already seen in Skylab EUV observations (see, e.g. Foukal 1975) and SOHO CDS (Del Zanna & Mason 2003), although with a much lower resolution (5'').

The Fe XII and Fe XIV predicted images show a more diffuse, unresolved emission. The AIA and Hi-C 193 Å images show both the resolved cool emission (mainly due to Fe VIII and Fe IX) and the unresolved 1–2 MK emission. It is interesting to note that there are very little differences in the AIA and Hi-C images, despite the factor of five higher resolution of Hi-C. This indicates that the hotter emission is not resolved even at the Hi-C resolution, while the cooler emission is already resolved at the AIA resolution. This is not an instrumental effect, because the low-lying moss emission is clearly much better resolved by Hi-C (cf. the lower right corner of the images in Fig. 13).

All the other coronal loops in the Hi-C FOV show similar features. As in the other ARs we have analysed, there are often locations where cooler and hotter loops are intermingled.

Fig. 14 shows region R2, where six cool loops appear as nearly resolved by AIA. Again, there is little difference between the Hi-C and AIA 193 Å images. But in this case there is no clear anti-correlation between the locations of the cooler loops and the unresolved higher-temperature emission.

Fig. 15 shows region R3, on a very bright leg of a coronal loop. It is quite clear that the strong cool emission is comprised of at least three loops, visible in Fe VIII and Fe IX (AIA 171 Å). It is also quite clear that some fainter, different loops are present in the 1–2 MK range, as seen in Fe XII. These loops are not co-spatial with the cooler ones and are visible in both the AIA and Hi-C 193 Å images. In this case, these hotter loops appear as slightly better resolved in the Hi-C 193 Å images.

10. Summary and conclusions

It is clear that significant advances in our understanding of the thermal structure in active regions can be achieved by combining high-resolution imaging with spectroscopy and atomic data. We have analysed several active regions, but showed here only a few selected examples. Clearly, not all active regions are the same, however several new important results have emerged from our small sample:

1) we have verified that our new Hinode EIS radiometric calibration is in very good agreement with the SDO AIA pre-flight and EVE in-flight calibrations. This gives us confidence in the

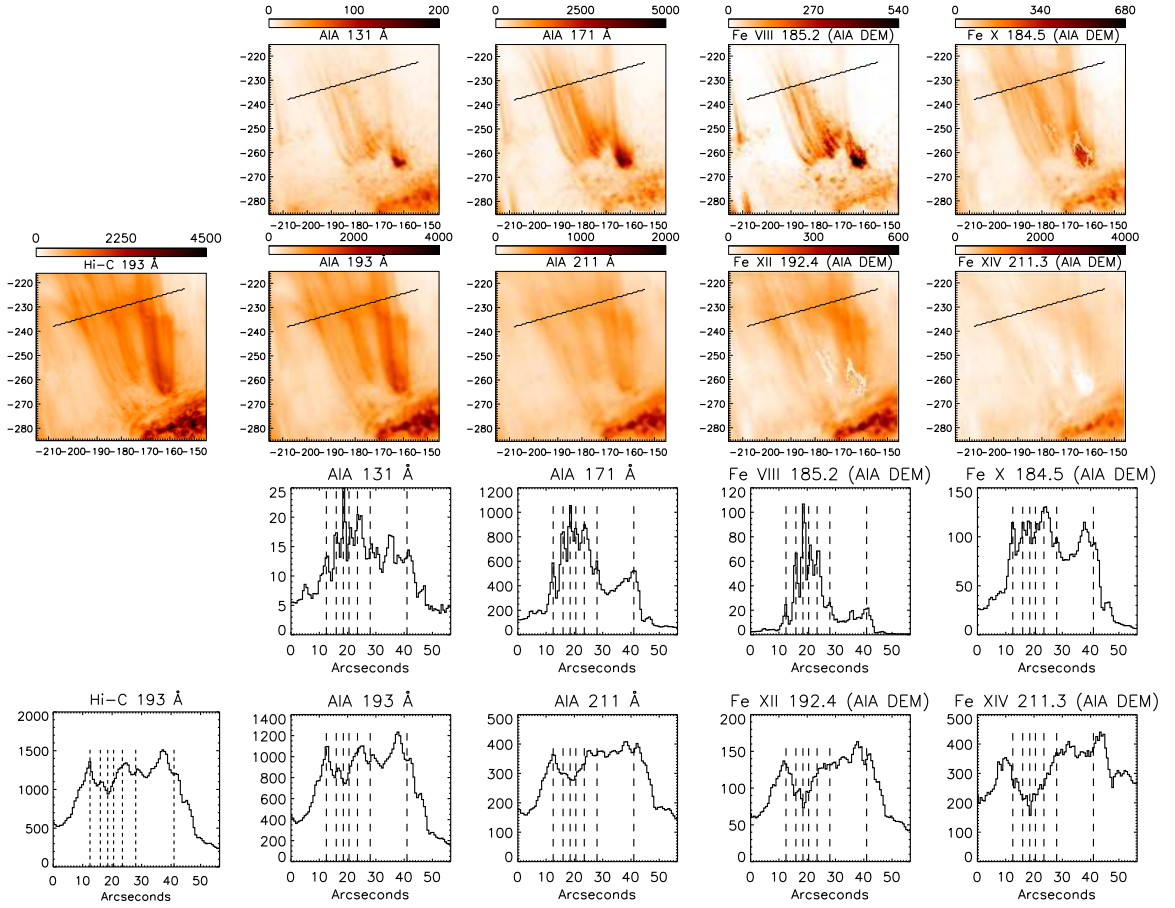


Fig. 13. Above: the region R1, as observed by Hi-C, by AIA, and as predicted from the AIA *DEM* modelling. Below: profiles (averages over $1.8''$ in DN/s) along the segment indicated in the images. The position of a few cool loops is noted.

present and future results. Some discrepancies are present, but they are within the combined uncertainties of the absolute calibrations of the AIA, EVE and EIS instruments (about 25% each).

2) We have verified that the atomic data we have provided for the EIS and AIA instruments are now accurate and sufficiently complete, with the exception of the AIA 211 Å band.

3) We have confirmed the multi-thermal nature of the AIA EUV bands, also in quiescent AR observations. In the cores of ARs, only the 171 and 335 Å bands are relatively isothermal. In the 193 and 211 Å bands, only about half of the counts are due to Fe XII and Fe XIV. We have provided ways to use the AIA data to roughly estimate the various contributions. Significant continuum emission is present in the 94 Å and, especially, the 131 Å band, confirming our previous suggestion (O'Dwyer et al. 2010). There is evidence that coronal abundances are more appropriate than the photospheric ones for the cores of active regions, but if the latter were the correct ones, the continuum contribution to the 131 Å band could be as high as 50%. The discussion of the various contributions to the AIA bands, based on the EIS spectra and the atomic data, was key to the proper use of the AIA data for *DEM* modelling.

4) After two years of atomic data calculations and benchmark, we have finally obtained a good agreement between observations and predictions for the AIA 94 Å band. We confirm our (Del Zanna 2012a) suggestion that the newly identified Fe XIV 93.61 Å line is the dominant contribution to the 94 Å band in any AR observation, whenever Fe XVIII is not present. Studies of hot AR emission should take this into account.

5) We have shown the difficulty in estimating the Fe XVIII contribution directly with the *DEM* modelling (first method). A better way is to subtract the various contributions to the observed count rates in the 94 Å band (second method). We used EIS spectroscopy and atomic data to provide a way to obtain rough estimates of the Fe XVIII emission in the 94 Å band using only AIA data (third method). We showed that Fe XVIII emission is sometimes present in the cores of ARs, although this is often caused by 3 MK and not by 7 MK plasma. The same applies to Fe XVII emission. These new results should be taken into account when using the 94 Å band to diagnose the presence of hot plasma. We showed the various complexities related to the calcium lines, and the surprising fact that Ca XVII is often not observed when Fe XVIII is present.

6) We provided what are to our knowledge the first consistent measurements of electron densities and elemental abundances in a sample of 3 MK loops. They are significantly different than previous measurements. Low-FIP elements such as iron are enhanced by about a factor of 3, compared to the high-FIP elements. We performed several *DEM* analyses on a few ARs cores and obtained similar results. The path lengths obtained from the densities and emission measures indicate that the abundances of the low-FIP elements are enhanced, compared to the photospheric values, by at least a factor of 3, which implies spectroscopic filling factors close to 1. If we assume that this is a real measure of how the plasma is distributed along the line of sight (see Judge 2000 for caveats), this suggests that we are still far from resolving the substructures in the 3 MK loops.

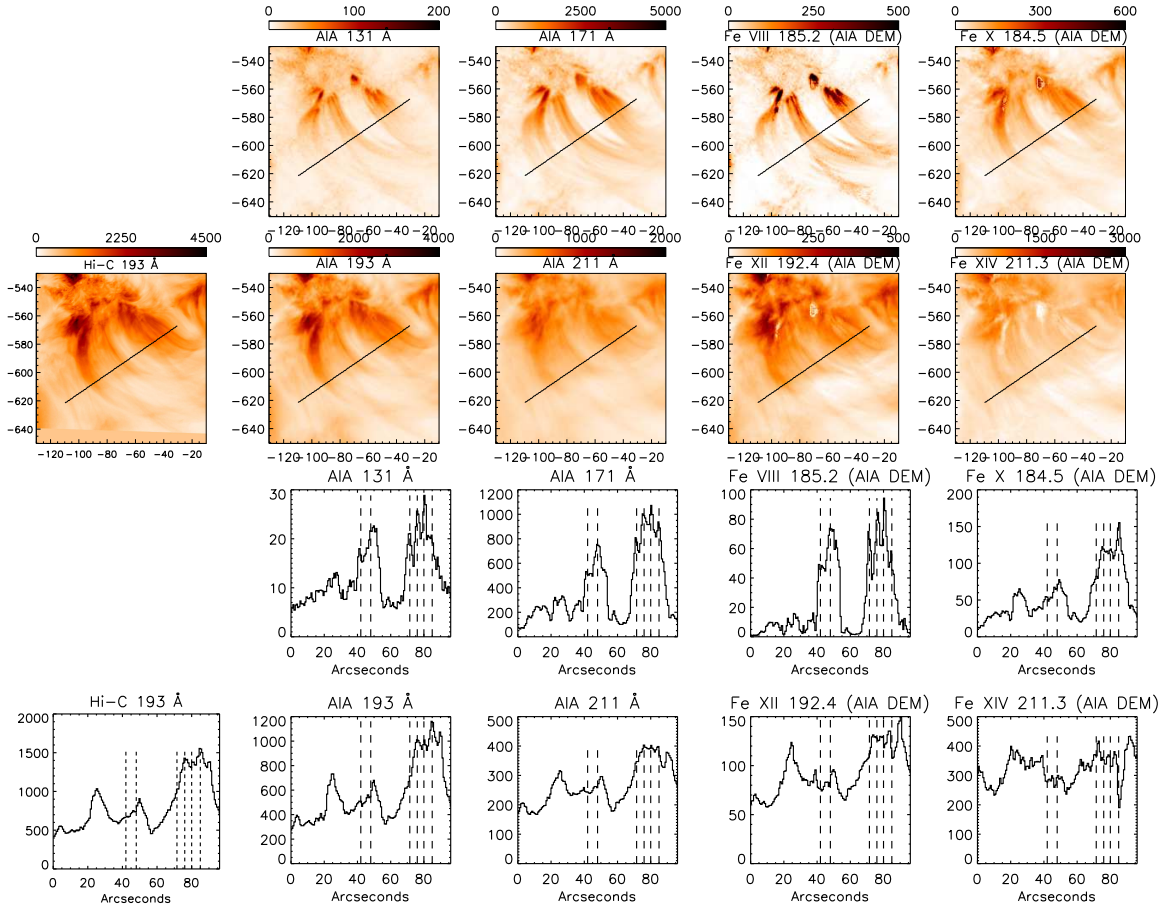


Fig. 14. Above: the region R2, as observed Hi-C, by AIA, and as predicted from the AIA DEM modelling. Below: profiles (averages over $1.8''$ in DN/s) along the segment indicated in the images. The position of a few cool loops is noted.

7) We have presented a new *DEM* method customised for analysing AIA observations. The comparison with direct spectroscopic observations from Hinode EIS is truly remarkable, providing confidence in using AIA for studying the thermal structure of ARs using the present method, atomic data, and caveats.

8) We often found, from the *DEM* modelling at the AIA resolution ($1''$), that loops emitting at different temperatures are not co-spatial, as we also saw at much lower resolution with SOHO CDS (Del Zanna et al. 2006). To undoubtedly identify the spatial location of all the loop structures visible along the line of sight would probably require stereoscopic isothermal measurements though.

9) The AIA *DEM* modelling at $1''$ resolution indicates that warm loops (up to 1 MK) are nearly resolved, and have very little background/foreground emission. It also indicates that the emission at 1.5–2.5 MK is progressively more unresolved, and that it increases towards the core of ARs. This unresolved emission was found to comprise about 80% of the emission in TRACE and SOHO/CDS observations (Del Zanna & Mason 2003). The AIA *DEM* modelling suggests much lower values. We recall that this reduction is mainly due to the lower stray light of the AIA telescopes (compared to TRACE) and to their higher spatial resolution (compared to CDS and EIS).

The observations of the 2 MK unresolved emission in an active region (Del Zanna 2012b) showed increases by about a factor of 2–2.5 of the iron abundance, compared to say argon. The absolute values (i.e. relative to hydrogen) are not known, but if we assume, as in the case of the 3 MK loops, that it is the abso-

lute iron abundance to increase, it follows that a factor of 2–3 enhancement in the observed emission is just due to this abundance variation and is not due to increased densities or temperatures.

In summary, the picture that is emerging is that of variable abundances within the different features in active regions. This is an important issue which complicates the interpretation of the observations, and that deserves further analysis.

10) We found very little differences in the AIA and Hi-C 193 Å images of coronal structures, despite the factor of five higher resolution of Hi-C. The AIA *DEM* modelling indicates that there is cool emission (mainly due to Fe VIII and Fe IX) in the Hi-C images which appears to be already nearly resolved at the AIA resolution. The Hi-C images also contain hotter (1.5–2.5 MK) emission which in most of the places appears as unresolved as in the AIA images. This suggests that fundamental sub-structuring is well below the $0.25''$ resolution at these temperatures.

Understanding the multi-thermal structure of active regions is clearly a complex issue. One key issue that is currently debated in the literature regards the temporal evolution of the various structures. For example, are some of the warm loops hot 3 MK loops observed while they are cooling? This is an even more complex issue that will be partly addressed in a separate publication, although it is really at the limits of current instrumentation.

Our SOHO CDS observations showed that the majority of the observed warm loops are long-lasting, i.e. do not appear to cool down or heat up significantly over timescales of hours (Del Zanna et al. 2006; Cirtain et al. 2007). Our results are in

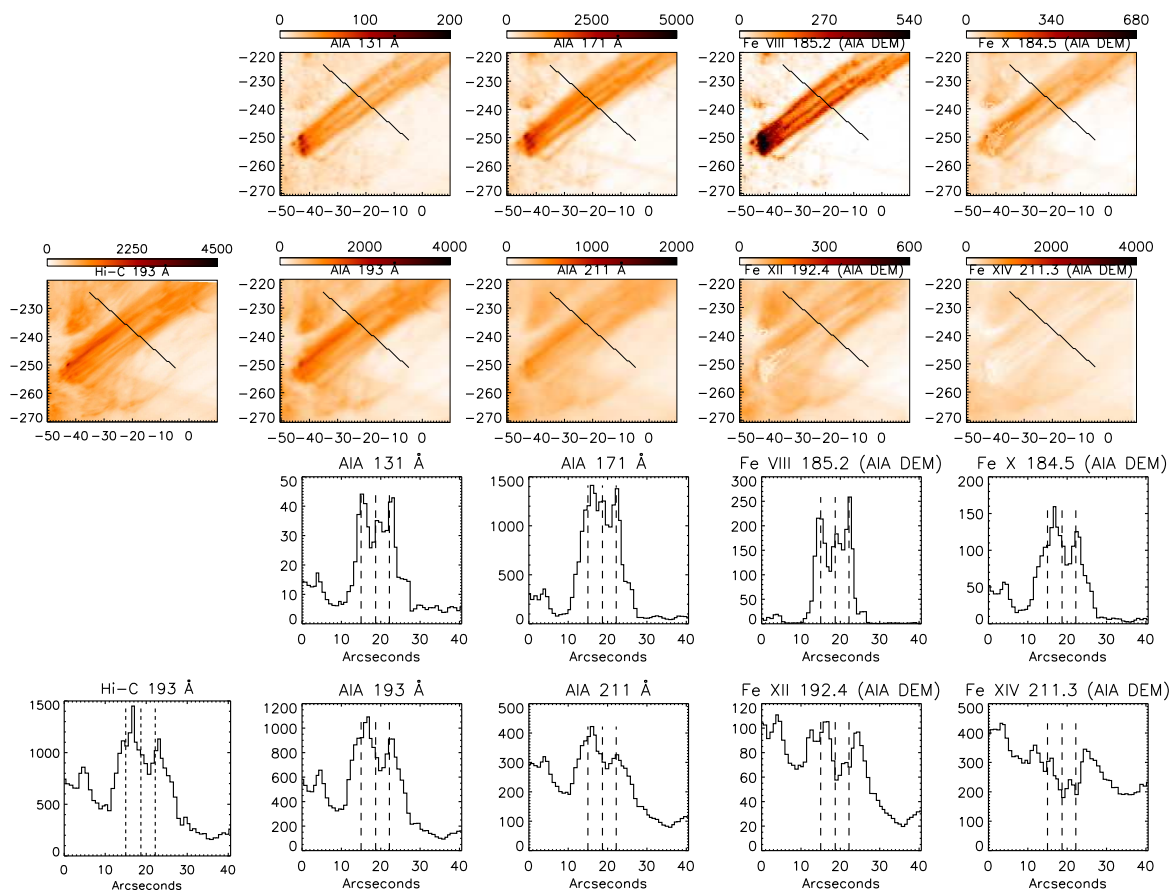


Fig. 15. Above: the region R3, as observed Hi-C, by AIA, and as predicted from the AIA DEM modelling. Below: profiles (averages over $1.8''$ in DN/s) along the segment indicated in the images. The position of a few cool loops is noted.

contrast to those obtained by Warren et al. (2003); Ugarte-Urra et al. (2006, 2009); Viall & Klimchuk (2011), although they do not necessarily imply steady heating for the single structures. Indeed the higher SDO AIA resolution shows clear variations on timescales of minutes. The quasi-steady temperature structure of loops reflects the topology of the photospheric magnetic field which appears relatively steady over timescales of hours. This does not mean that the coronal magnetic field is steady. For example, according to our model of AR evolution (Del Zanna et al. 2011a), interchange reconnection should occur between the 3 MK loops and the surrounding magnetic field, producing as a byproduct the formation of warm loops. This probably occurs on very short timescales and small spatial scales, well below the resolution of current instrumentation. It is however a long-lasting (hours to days) quasi-steady process.

The other key issue that is currently debated regards the sub-resolution structures. We believe that significant advances in the future can only be obtained with high-resolution spectroscopy (e.g. the proposed Solar-C spectrometer, see Teriaca et al. 2012) and isothermal imaging. The current imaging at 1 (AIA) and $0.25''$ (Hi-C) resolution indicates, however, that loops at temperatures above 1.5 MK are still far from being resolved. We suggest that future high-resolution imaging with Hi-C is done with isothermal bands, possibly covering the 0.7–1.5 MK temperature range.

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