Mathematical Tripos Part III Lent Term 2020
Black Holes: Examples Sheet 3 H.S. Reall

1. Consider the Reissner-Nordstrom solution with $M > e$ using advanced Eddington-Finkelstein coordinates. (a) Determine the Finkelstein diagram (i.e. show ingoing and outgoing radial null geodesics in a plot of $t_* = v - r$ against r). (b) Show that r decreases along any causal curve in the region $r_- < r < r_+$.

2. (a) Prove that, if a vector field ξ preserves the Maxwell field (i.e. $\mathcal{L}_\xi F = 0$) then locally there exists a scalar potential Φ such that $i_\xi F = d\Phi$. (*Hint:* Q2 of examples sheet 1.)

   (b) The equation of motion of a particle of charge q and 4-velocity U^a is $U^b \nabla_b U^a = (q/m) F^a{}_b U^b$. Let ξ be a Killing vector field that preserves the Maxwell field. Show that $\xi \cdot U - (q/m)\Phi$ is conserved along the particle's worldline.

   (c) Deduce that, for a particle of mass m moving in the equatorial plane ($\theta = \pi/2$) of a Reissner-Nordstrom black hole (with $Q > 0$, $P = 0$), the quantities $\mathcal{E} = (\Delta/r^2) dt/d\tau + qQ/(mr)$, and $h = r^2 d\phi/d\tau$ are constant (τ is proper time). Hence show that the radial motion is determined by the equation
   $$\left(\frac{dr}{d\tau}\right)^2 + \frac{\Delta(r)}{r^2}\left(1 + \frac{h^2}{r^2}\right) = \left(\mathcal{E} - \frac{qQ}{mr}\right)^2 .$$

   (d) What is the physical interpretation of the case $q/m = Q/M = 1$, $\mathcal{E} = 1$, $h = 0$?

   (e) *The Penrose process.* A particle P_1 falls from $r = \infty$ towards the black hole. Just before it crosses the event horizon, it decays into two other particles P_2 and P_3 where P_2 has charge $q < 0$. The decay happens such that P_2 initially has $dr/d\tau \approx 0$. P_2 subsequently falls into the black hole and P_3 escapes to $r = \infty$. Let $E_i \equiv m_i \mathcal{E}_i$ denote the energy of P_i (which has mass m_i). Show that $E_1 > 0$ and $E_2 < 0$. Hence, by energy conservation, $E_3 > E_1$, i.e., the particle returning to infinity has more energy than the initial particle! This is consistent because P_2 has carried negative energy into the black hole. Hence energy (and charge) are extracted from the black hole in this process.

3. Let E and h be the energy and and angular momentum per unit mass of a zero charge free particle moving in the equatorial plane ($\theta = \pi/2$) of a Kerr-Newman black hole. Show that the particle's Boyer-Lindquist radial coordinate r satisfies
   $$\frac{1}{2}\left(\frac{dr}{d\tau}\right)^2 + V(r) = \frac{1}{2}E^2 ,$$
   where τ is an affine parameter, and the effective potential V is given by
   $$V(r) = \frac{1}{2}\left(1 - \frac{2M}{r} + \frac{e^2}{r^2}\right)\left(\sigma + \frac{h^2}{r^2}\right) + \frac{aEh}{r^3}\left(2M - \frac{e^2}{r}\right) + \frac{a^2}{2r^2}\left[\sigma - E^2\left(1 + \frac{2M}{r} - \frac{e^2}{r^2}\right)\right] .$$
   where $\sigma = 1$ for a massive particle and $\sigma = 0$ for a massless particle. In the Kerr case ($e = 0$), obtain an expression for the radius of orbits of constant r. Deduce the radius of ISCO.

4. Obtain the Kerr-Newman solution in Kerr coordinates. When is there a regular event horizon? Show that the area of the event horizon of a Kerr-Newman black hole is $A = 8\pi(M^2 - e^2/2 + \sqrt{M^4 - e^2 M^2 - J^2})$.

5. Let E denote the maximum energy that can be extracted from a Kerr black hole in the Penrose process. The *efficiency* of this process is $\eta \equiv E/M$ where M is the initial mass of the black hole. What is the largest possible value of η?

6. In the Kerr geometry, consider two spacelike surfaces Σ, Σ' which both extend from i^0 to $\mathcal{H}^+$ with Σ' lying entirely to the future of Σ. Let H and H' denote the intersections of Σ and Σ' with $\mathcal{H}^+$. Let $\mathcal{N}$ denote the portion of $\mathcal{H}^+$ from H to H'. Let $J_a = -T_{ab}k^b$ be the conserved energy-momentum 4-vector.

   (a) Show that
   $$E(\Sigma') - E(\Sigma) = \int_{\mathcal{N}} \star J, \tag{1}$$
   where $E(\Sigma) \equiv -\int_\Sigma \star J$ is the total energy of matter fields on Σ, and similarly for $E(\Sigma')$. What is the physical interpretation of this formula?

   (b) Explain why the orientation of $\mathcal{N}$ used in this formula is given by $dv \wedge d\theta \wedge d\chi$ in Kerr coordinates (the orientation of spacetime is given by $dv \wedge dr \wedge d\theta \wedge d\chi$).

   (c) Show that $(\star J)_{v\theta\chi} = (r_+^2 + a^2)\sin\theta \xi^a J_a$.

   (d) Assume that matter obeys the dominant energy condition. Explain why $E(\Sigma') \le E(\Sigma)$ for a Schwarzschild black hole (i.e. $a = 0$) but why this is not necessarily true for a Kerr black hole.

   (e) Now take the matter to be a massless real scalar field, with energy-momentum tensor $T_{ab} = \partial_a\Phi\partial_b\Phi - (1/2)g_{ab}(\partial\Phi)^2$. Consider a mode of this field with frequency ω and azimuthal quantum number ν, i.e., $\Phi = \mathrm{Re}[\Phi_0(r,\theta)\exp(-i\omega v + i\nu\chi)]$. Show that the RHS of equation (1) is positive for $0 < \omega < \nu\Omega_H$. (Note that $\xi \cdot k = 0$ on $\mathcal{H}^+$ because $\mathcal{H}^+$ must be invariant under an isometry hence any Killing field must be tangent to $\mathcal{H}^+$.)

   This example shows that energy can be extracted from a black hole by scattering waves off it. This is called *superradiant scattering.*

7. (a) Let (M, g) be a stationary vacuum spacetime containing an hypersurface Σ such that the initial data induced on Σ is geodesically complete and asymptotically flat with 1 end. Prove that the Komar mass must vanish and hence, by the positive energy theorem, that the spacetime must be flat. (This is a version of *Lichnerowicz's theorem* which excludes the existence of gravitational solitons, i.e., stationary configurations of the gravitational field that are not black holes.)

8. (a) Calculate the ADM mass of the Reissner-Nordstrom solution.

   (b) Calculate the electric and magnetic charges, the Komar mass, and the Komar angular momentum of the Kerr-Newman solution.